

# Northumbria Research Link

Citation: Syntelis, P. and Antolin, Patrick (2019) Kelvin-Helmholtz Instability and Alfvénic Vortex Shedding in Solar Eruptions. *The Astrophysical Journal Letters*, 884 (1). L4. ISSN 2041-8205

Published by: IOP Publishing

URL: <https://doi.org/10.3847/2041-8213/ab44ab> <<https://doi.org/10.3847/2041-8213/ab44ab>>

This version was downloaded from Northumbria Research Link:  
<http://nrl.northumbria.ac.uk/id/eprint/41420/>

Northumbria University has developed Northumbria Research Link (NRL) to enable users to access the University's research output. Copyright © and moral rights for items on NRL are retained by the individual author(s) and/or other copyright owners. Single copies of full items can be reproduced, displayed or performed, and given to third parties in any format or medium for personal research or study, educational, or not-for-profit purposes without prior permission or charge, provided the authors, title and full bibliographic details are given, as well as a hyperlink and/or URL to the original metadata page. The content must not be changed in any way. Full items must not be sold commercially in any format or medium without formal permission of the copyright holder. The full policy is available online: <http://nrl.northumbria.ac.uk/policies.html>

This document may differ from the final, published version of the research and has been made available online in accordance with publisher policies. To read and/or cite from the published version of the research, please visit the publisher's website (a subscription may be required.)



**Northumbria  
University**  
NEWCASTLE



**University Library**



# Kelvin–Helmholtz Instability and Alfvénic Vortex Shedding in Solar Eruptions

P. Syntelis<sup>1</sup> and P. Antolin<sup>2,1</sup>

<sup>1</sup> St Andrews University, Mathematics Institute, St Andrews KY16 9SS, UK; [ps84@st-andrews.ac.uk](mailto:ps84@st-andrews.ac.uk)

<sup>2</sup> Department of Mathematics, Physics and Electrical Engineering, Northumbria University, Newcastle Upon Tyne NE1 8ST, UK

Received 2019 July 19; revised 2019 September 5; accepted 2019 September 12; published 2019 October 3

## Abstract

We report on a three-dimensional MHD numerical experiment of a small-scale coronal mass ejection (CME)-like eruption propagating through a nonmagnetized solar atmosphere. We find that the Kelvin–Helmholtz instability (KHI) develops at various but specific locations at the boundary layer between the erupting field and the background atmosphere, depending on the relative angle between the velocity and magnetic field. KHI develops at the front and at two of the four sides of the eruption. KHI is suppressed at the other two sides of the eruption. We also find the development of Alfvénic vortex shedding flows at the wake of the developing CME due to the 3D geometry of the field. Forward modeling reveals that the observational detectability of the KHI in solar eruptions is confined to a narrow  $\approx 10^\circ$  range when observing off-limb, and therefore its occurrence could be underestimated due to projection effects. The new findings can have significant implications for observations, for heating, and for particle acceleration by turbulence from flow-driven instabilities associated with solar eruptions of all scales.

*Unified Astronomy Thesaurus concepts:* Magnetohydrodynamics (1964); Solar activity (1475); Solar active regions (1974); Solar atmosphere (1477); Solar atmospheric motions (1478); Solar corona (1483); Solar coronal mass ejections (310); Solar magnetic flux emergence (2000); Solar magnetic fields (1503); Magnetohydrodynamical simulations (1966)

## 1. Introduction

The Kelvin–Helmholtz instability (KHI) can occur at the boundary layer between two moving fluids, when the velocity shear between them is such that it overcomes the surface tension of the boundary layer. KHI is a very common physical process and has been associated with many astrophysical systems, such as, for example, planetary magnetospheres (Hasegawa et al. 2004; Masters et al. 2010; Slavin et al. 2010), astrophysical jets (e.g., Ferrari et al. 1981; Bodo et al. 1994), and disks (e.g., Balbus 2003; Johansen et al. 2006). In solar-related phenomena, KHI has been observed at the flanks of coronal mass ejections (CMEs; Foulon et al. 2011, 2013; Möstl et al. 2013), in the dimming regions associated with solar eruptions (Ofman & Thompson 2011), at the sides of solar jets (Kuridze et al. 2016; Li et al. 2018, 2019), and at coronal streamers (Feng et al. 2013). KHI can play a very important role in the development of turbulence, the mixing and heating of plasma (e.g., Heyvaerts & Priest 1983; Terradas et al. 2008; Antolin et al. 2016; Fang et al. 2016), and particle acceleration (Leroy et al. 2018).

The threshold and growth rate of KHI at the boundary layer between two horizontal magnetized flows was derived analytically by Chandrasekhar (1961). The theory has been since extended for flux tubes in the solar convection zone (Tsinganos 1980), for boundaries between a twisted flux tube and a twisted or a horizontal external field (Zaqarashvili et al. 2014a, 2014b, 2015; Zhelyazkov et al. 2015a, 2015b), and also for the effects of partial ionization (Martínez-Gómez et al. 2015). A major point from the theory is that the onset and the growth rate of KHI depend on the angle between the shear flow and the magnetic field at the boundary layer. KHI can develop easier in a shear flow perpendicular to a magnetic field, whereas KHI can be suppressed in a shear flow parallel or antiparallel to a magnetic field (e.g., Hillier et al. 2019).

KHI in the context of solar eruptions has been studied using models that describe shear flows between two vertical regions, mimicking locally the boundary layer of the flank of a CME (Ofman & Thompson 2011; Nykyri & Foulon 2013; Möstl et al. 2013; Gómez et al. 2016). In another approach, KHI has been studied by assuming magnetic cylinders moving through a model atmosphere, mimicking the ejection of a flux rope (Pagano et al. 2007). Such models have the advantage of being easy to implement. As such, they are useful to perform controlled parametric studies. However, the magnetic field of a full 3D CME eruption is significantly more complicated than such idealized configurations (e.g., Syntelis et al. 2017). In this study, we report on the first 3D MHD simulation of a solar eruption showing the development of KHI at the boundary layer between the erupting field and the background atmosphere. We show that the KHI only develops at specific boundaries where the shear flow is mostly perpendicular to the local magnetic field, with important implications for assessing its occurrence as well as the generation of turbulence during CME propagation.

## 2. Model

The model used in this Letter is described in detail in Syntelis et al. (2017, 2019). We use Lare3d (Arber et al. 2001) to solve the 3D time-dependent, resistive, and compressible MHD equations in Cartesian geometry. The equations, resistivity form, and normalization can be found in Syntelis et al. (2017).

The numerical domain has a physical size of  $153^3$  Mm ( $1000^3$  grid points). The boundary conditions are periodic in the  $y$ -direction and open in the  $x$ -direction. In the  $z$ -direction the boundaries are open (closed) for the top (bottom) of the domain.

The domain consists of a solar interior and a solar atmosphere. The interior is a convective stable and adiabatically stratified layer between  $-7.2$  Mm  $\leq z < 0$  Mm. The atmospheric temperature

follows a hyperbolic tangent profile for the temperature, mimicking an isothermal photospheric–chromospheric layer ( $0 \text{ Mm} \leq z < 1.8 \text{ Mm}$ ), a transition region ( $1.8 \text{ Mm} \leq z < 3.2 \text{ Mm}$ ), and an isothermal corona ( $3.2 \text{ Mm} \leq z < 145.8 \text{ Mm}$ ). The atmospheric density and pressure are derived by assuming hydrostatic equilibrium.

A horizontal flux tube is positioned inside the solar interior, along the  $y$ -direction, at  $z = -2.1 \text{ Mm}$ . The flux tube's magnetic field in cylindrical coordinates is

$$B_y = B_0 \exp(-r^2/R^2), \quad (1)$$

$$B_\phi = \alpha r B_y, \quad (2)$$

where  $R = 450 \text{ km}$  is the tube's radius,  $r$  is the radial distance from the tube axis,  $\alpha = 0.0023 \text{ km}^{-1}$  is a proxy of the twist of the tube, and  $B_0 = 2400 \text{ G}$  is the magnetic field strength at its center. The flux tube is initially in pressure equilibrium. To initiate the emergence of the flux tube a density deficit is imposed along the tube's axis (e.g., Fan 2001):

$$\Delta\rho = \frac{p_t(r)}{p(z)} \rho(z) \exp(-y^2/\lambda^2), \quad (3)$$

where  $p$  is the external pressure,  $p_t$  is the total pressure within the flux tube, and  $\lambda = 0.9 \text{ Mm}$  is the length scale of the buoyant part of the flux tube.

The formation of KHI depends on the relative angle of the shear flow and the direction of the magnetic field. We assume a nonmagnetized background atmosphere to assess the locations at the surface of the eruption where KHI can develop, solely due to the direction of the erupting field, without the influence of a preferred direction imposed by an ambient field. The inclusion of an ambient field can potentially affect the locations where KHI develops and make KHI less prominent (e.g., Pagano et al. 2007; see Section 4).

The real viscosity in the solar corona is estimated to be many orders of magnitude smaller than the numerical viscosity in our simulation. Unrealistically high viscosities can suppress the formation of KHI by significantly reducing the velocity shear between the two interfaces (e.g., Howson et al. 2017a; Antolin & Van Doorsselaere 2019). It is possible that with higher spatial resolution and correspondingly lower numerical viscosity, additional KHI could be obtained in our model. In this study, we have not examined the effects of resolution and numerical viscosity to the development of KHI as it would be very demanding computationally.

### 3. Results

Figure 1 shows the time evolution of the temperature (first row) and  $x$ -component of the vorticity (second row) before and during the eruption. The top left panel shows the temperature distribution at the vertical midplane, inside which a low-lying flux rope is located. This flux rope has been formed by the low-lying gradual tether-cutting reconnection of sheared lines (details in Syntelis et al. 2017, 2019). During the eruption (middle column), the erupting field (structure between  $30 < z < 110 \text{ Mm}$ , (b)) develops a hot edge layer (outer part of the field). We will refer to the outermost part of the edge as the boundary layer between the magnetized eruption and the nonmagnetized atmosphere. The temperature at the very edge of the boundary layer is cooler than that of the background atmosphere, and cooler than that of the hot edge part of the eruption. This is due to the adiabatic expansion of

this part of the boundary layer, as the field strength there decreases to zero.

During the eruption, the erupting field displaces the atmospheric material above it. The displaced material is pushed away from the uppermost part of the erupting field and then down beside the sides of the eruption. Therefore, velocity shear develops at the boundary layer, both by the upward velocity of the erupting field and by the sideways and downward velocity of the displaced material. This velocity shear can be seen as regions of high vorticity (Figures 1(d)–(f)).

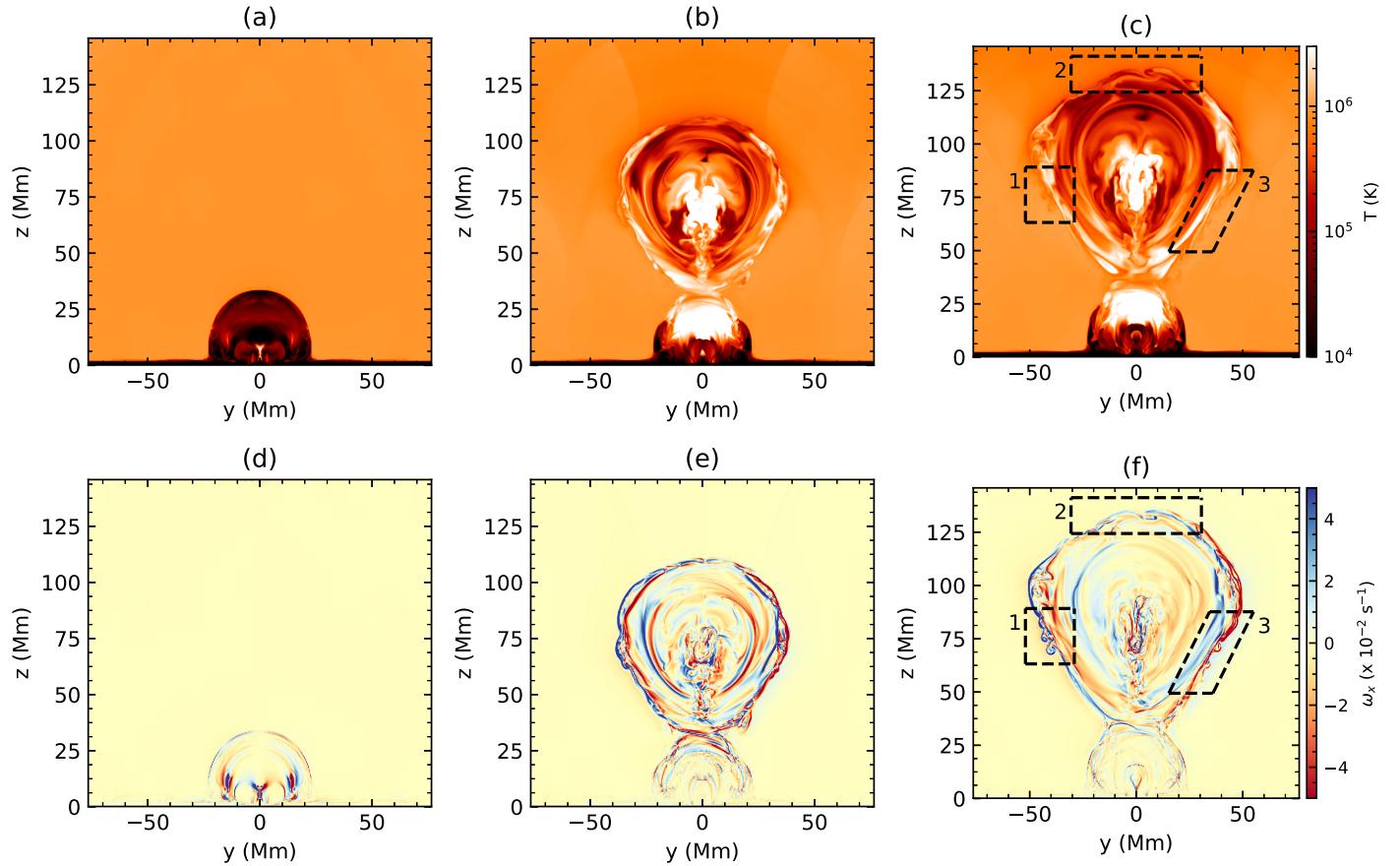
KHI develops at three locations at the boundary layer (Figures 1(c), (f)). Regions 1 and 3 are the left and right flank of the eruption when observing the structure along the  $+x$ -direction. Region 2 is the upper part of the erupting field. There, KHI vortices are larger in size.

Region 1 is examined in more detail in Figure 2. The temperature (a) and density (b) of the wave-like features take values between that of the background atmosphere and that of the boundary layer prior to the development of the waves (Figure 1(b)). However, KHI has not yet developed enough to mix the hot edge of the eruption (right of the wave-like features) with the atmosphere (left of the wave-like features). This is similar to the initial stage of the mixing found in simulations of KHI in coronal loops due to transverse MHD waves (e.g., Antolin et al. 2014; Howson et al. 2017a, 2017b; Karampelas et al. 2017). The plasma  $\beta$  inside the vortices, although mostly high (blue/white, panel (c)), increases locally at the fronts of the waves (white/red), indicating the local compression of the plasma due to the instability.

Panel (d) shows the  $\omega_x$  component of the vorticity. The wave-like features have mostly positive  $\omega_x$  (blue), indicating a consistent counterclockwise rotation in the  $yz$ -plane. This rotation is also reflected in the streamlines of the planar velocity field (green lines, (e)). The negative values of  $\omega_x$  (red) near the base of the wave-like features are associated mostly with shearing and not with an actual rotation.

Inside the central part of two lower wave-like features ( $z \approx 70\text{--}75 \text{ Mm}$ ),  $\omega_x$  changes sign. This is due to the off-plane motion of the vortices, indicating that the KHI vortices have a 3D structure. To examine that off-plane motion, we plot the ratio of the magnitude of the planar velocity over the magnitude of the total velocity (panel (e)). White indicates a planar flow, and dark blue indicates a flow away from the plane. Similarly, in panel (f) we plot the planar magnetic field strength over the total magnetic field strength. The velocity field shows a component away from the plane, contrary to the magnetic field, implying that the velocity shear is at an angle locally to the magnetic field, therefore favoring the development of KHI. Also, the KHI features develop internally a 3D structure where the velocity, magnetic, and vorticity fields have strong off-plane components.

We now examine the 3D structure of the whole boundary layer of the eruption (Figure 3). In panel (a), we trace magnetic field lines from the edges of the eruption (not the KHI regions). The eruption displaces the material above it and, moving it away from the apex of the eruption, toward every direction (green arrows). In some locations the flow is perpendicular to the magnetic field (favoring KHI), while in other locations the flow is parallel or antiparallel to the field (not favoring KHI). To visualize this, in panel (b) we plot the isosurface of the magnitude of vorticity colored by  $\hat{\mathbf{v}} \cdot \hat{\mathbf{B}}$ , where  $\hat{\mathbf{v}}$  and  $\hat{\mathbf{B}}$  are the unit vectors in the direction of the velocity and the magnetic



**Figure 1.** Time evolution of temperature (top row) and  $x$ -component of the vorticity (bottom row) at the  $yz$ -midplane. The first, second, and third columns correspond to  $t = 107.8, 128.6$ , and  $135.7$  minutes. The boxed regions in (c), (f) show regions where KHI has developed.

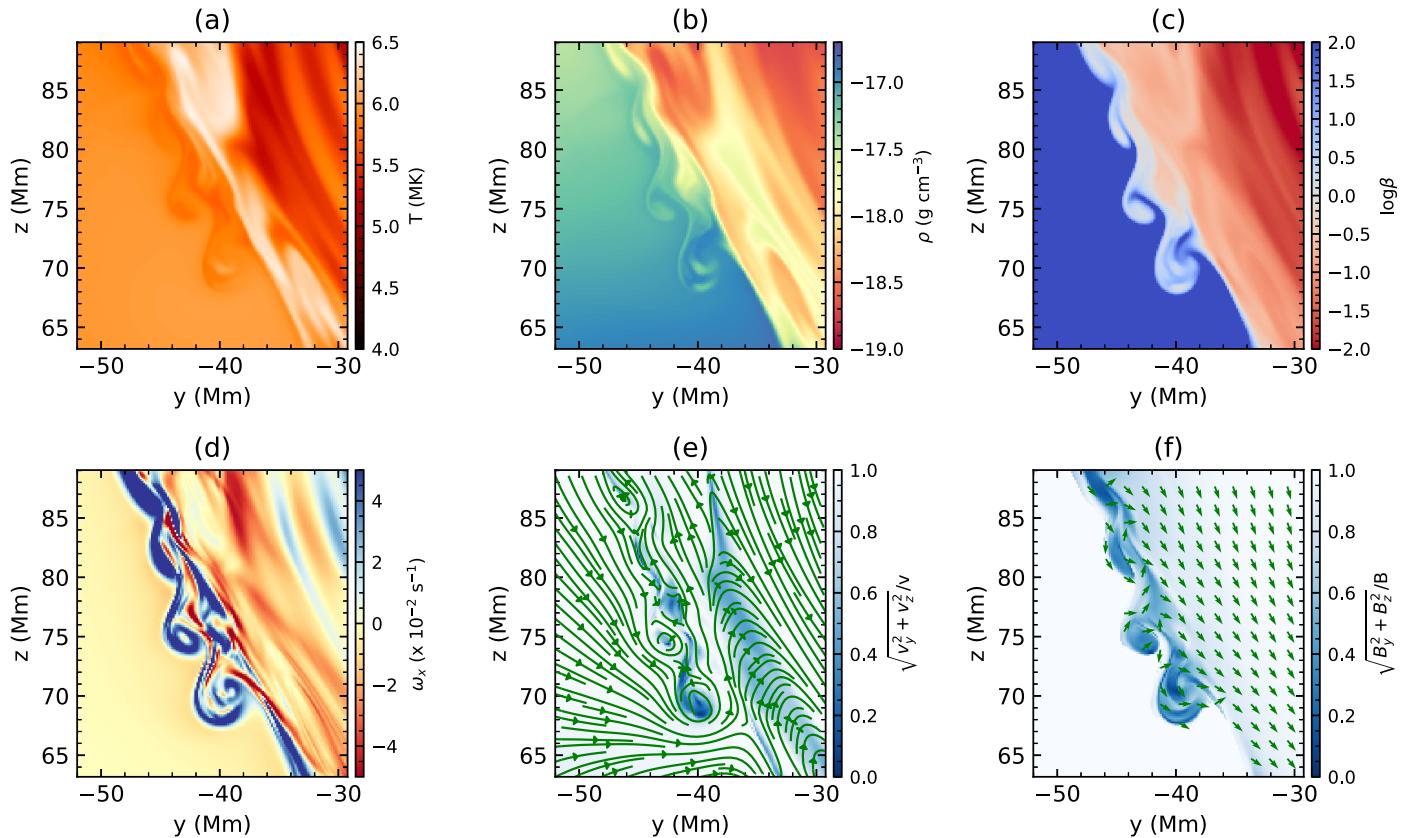
field. Green indicates the locations favoring KHI, while red and blue indicate the locations not favoring KHI. The 3D visualization reveals that a significant part of the surface of the eruption is covered with wave-like features of different sizes, thereby corrugating the surface. These features have maximum amplitude approximately along the dashed line, where  $\hat{v} \cdot \hat{B} \approx 0$ . The numbers 1, 2, 3 denote approximately the locations of the 2D boxes of Figure 1. Moving from the green regions toward the red (c) or blue (d) regions, the amplitudes of the wave-like features decrease and eventually become zero (arrows showing minimum KHI). As a result, at the two sides of the eruption associated with the red and blue regions, there are no signs of KHI. At the two other sides, the field geometry is such that the KHI develops. The complex 3D structure of KHI at various locations on the surface of the eruption implies that projection effects could be highly influencing the observational detection of these features.

To assess the projection effects, we produce synthetic observations of Fe IX along two lines of sight using the FoMo code (Van Doorsselaere et al. 2016; Figure 4), one being perpendicular to the plane of Figure 1 ( $x$ -axis,  $90^\circ$  angle; see the black arrow in Figure 3(b)). At that angle regions 2 and 3 show no sign of KHI (a). Panels (c)–(e) show the line intensity, Doppler velocity, and line width of region 1 (comparable to Figure 2). For simplicity, the spectral features have been obtained by fitting single Gaussians to the spectral profiles. However, the KHI introduces multiple spectral components, and the Doppler velocities shown here do not represent the extrema. These are represented by the broad line widths obtained from the single Gaussian fits, which are

mostly of nonthermal origin. The KHI features are visible in the intensity image; however, they are more prominent in the Doppler velocity and line width images. The second line of sight is approximately perpendicular to the maximum KHI line of Figure 3(b) ( $45^\circ$  angle, gray arrow). Region 2 now shows some undulating features, whereas regions 1 and 3 do not show signs of KHI. Panels (f) and (g) show a close-up of region 2 for the two angles. In panel (f) there is no indication of KHI. In panel (g), the undulating feature is present; however, it does not appear as KHI waves, as they are “smoothed” due to the integrating effect of the optically thin radiation across the complex region of the top of the eruption. Therefore, observations of such features are not likely to be associated with KHI. The Doppler velocity (h) and line width (i) at the  $45^\circ$  angle again produce clear signatures of these undulating features, however, without showing characteristic KHI crests. The complexity of region 2 can be witnessed in the line width, where the emission from plasma flowing at different directions results in very wide line profiles with high nonthermal broadening.

To further assess the projection effects, in Figure 5 we show the line intensity of region 2 for nine angles between  $10^\circ$  and  $90^\circ$ . The undulating features appear more clearly between  $40^\circ$  and  $50^\circ$  and become smoothed away from these angles. In general, for lower-resolution and noisier data, such features would be further degraded. Similar results are obtained for regions 1 and 3. Therefore, current observations might be underestimating the presence of KHI in solar eruptions due to line-of-sight projection and resolution effects.

Besides features associated with KHI, in Figures 3(c) and (d), we find other localized regions of vorticity. These correspond to



**Figure 2.** Close-up of region 1, shown in Figure 1. (a) Temperature, (b) density, (c) plasma  $\beta$ , and (d)  $x$ -component of the vorticity. (e) The magnitude of the planar velocity over the total velocity. Overplotted are streamlines of the planar velocity field. (f) The magnitude of the planar magnetic field over the total magnetic field. Overplotted is the planar magnetic field vector.

Alfvénic vortex shedding developing at two of the sides of the eruption (“vortex shedding” arrows). These regions are more clearly shown in Figure 3(e). In panel (f), we plot the magnetic field strength at a vertical slice through the vortex shedding regions of panel (e). At the two sides of panels (e), (f), the field geometry is such that the “legs” and the upper part of the eruption form two convex regions. The flow moving around the sides of the upper part of the eruption and toward the narrow point between the convex regions becomes turbulent. The drag forces deform the boundary layer, forming these characteristic thin and turbulent structures associated with vortex shedding (Bonet et al. 2008; Nakariakov et al. 2009; Gruszecki et al. 2010).

#### 4. Discussion

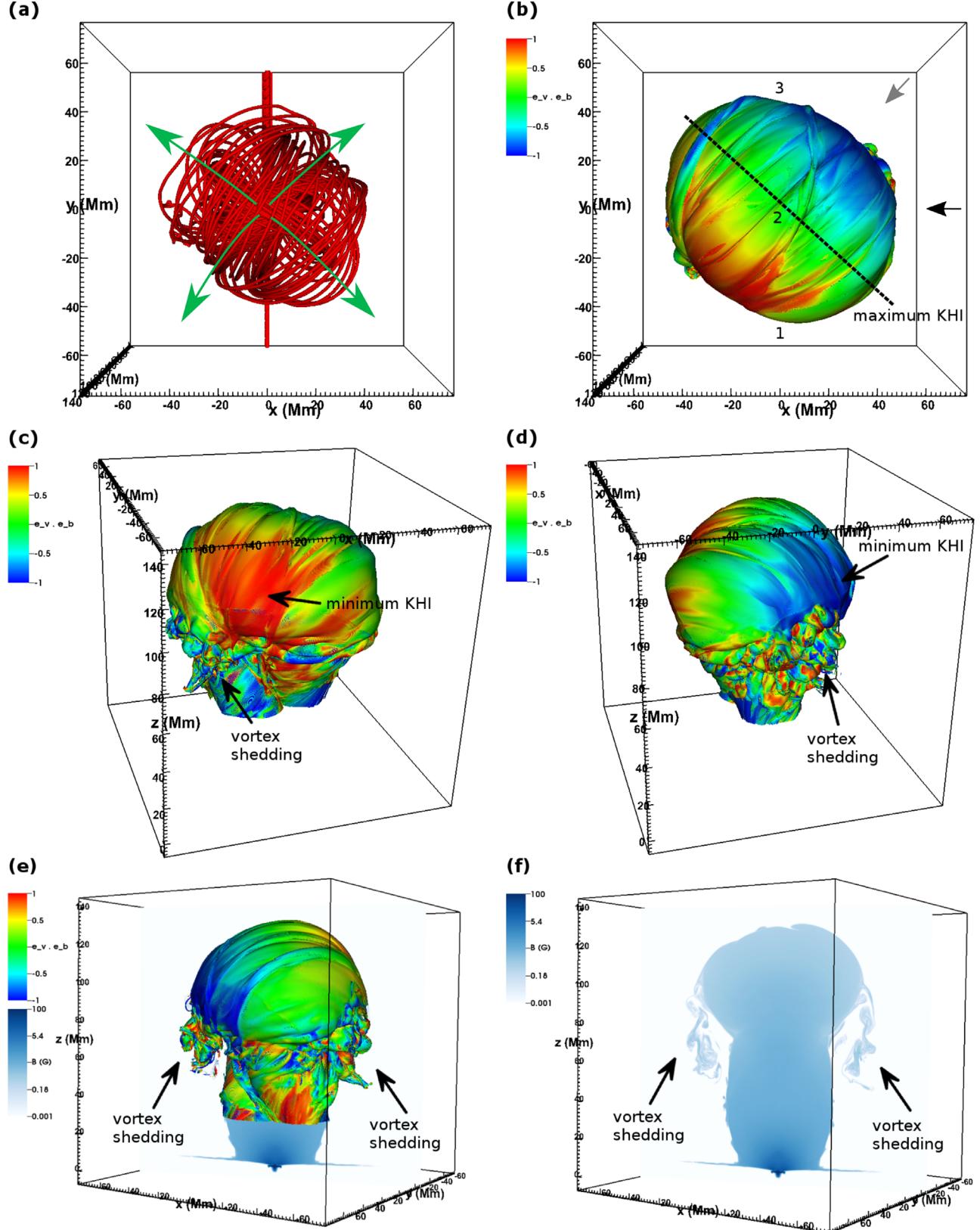
We have presented results on the formation of KHI and vortex shedding in a simulation of a small-scale solar eruption. The results demonstrate that using a realistic 3D model of a solar eruption is crucial to assess the formation and the effects of KHI and other flow-driven instabilities in the local development of turbulence, the heating of plasma, and any possible effects associated with such processes (e.g., small-scale reconnection, local acceleration of particles; Nykyri & Otto 2004; Fang et al. 2016; Leroy et al. 2018; Páez et al. 2019).

The reported simulations have been performed in a nonmagnetized atmosphere. Adding an ambient field would affect the formation of KHI (e.g., Pagano et al. 2007). Changes on the local plasma properties at the boundary layer (e.g., shear velocity, boundary layer width, relative magnetic field orientation, magnetic field strength, compressibility of the fluid, etc.)

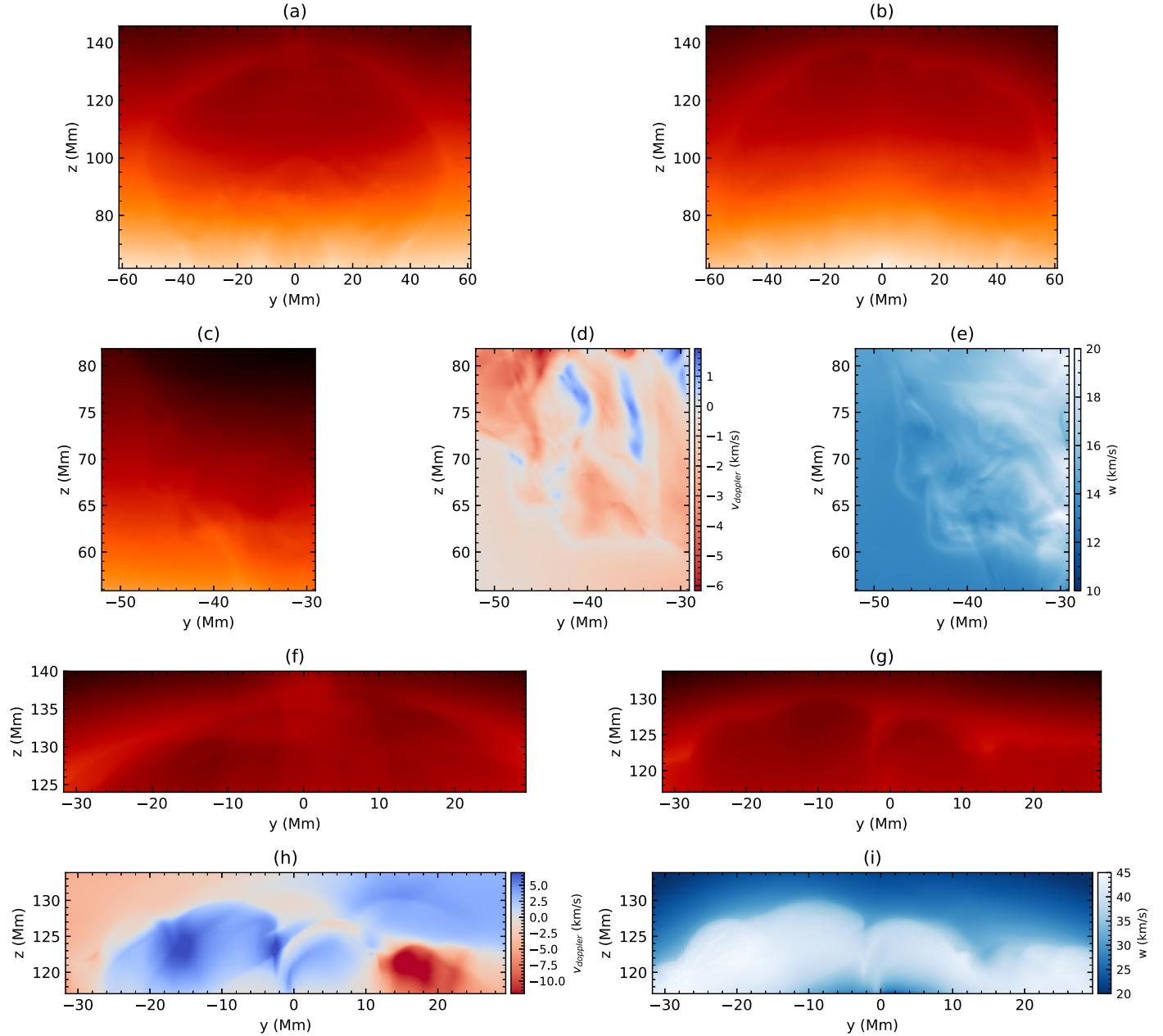
can lead to, e.g., changes in the growth rate of the instability and affect the sizes, shapes, and mixing of the vortices. Such effects can make the instability less prominent or even suppress it, depending on the specific local properties (e.g., Miura & Pritchett 1982; Ryu et al. 2000; Möstl et al. 2013; Nykyri & Foulon 2013; Tian & Chen 2016; Faganello & Califano 2017; Ma et al. 2017). Also, an ambient field would affect the motions of the displaced material, changing the angle of the shear and thus affecting the locations where KHI can develop. Möstl et al. (2013) showed that depending on the relative angles between the magnetic field and the velocity shear along both sides of a horizontal magnetized layer, the layer could become unstable at only one of its two sides. Given the complexity of the 3D magnetic field of a solar eruption (Figure 3(a)), the presence of an ambient field could potentially suppress KHI in one of the two unstable sides of the eruption, allowing KHI to develop only in one side of the eruption, similar to the observations of Foulon et al. (2011). However, our results strongly suggest that such a case can be easily attributed to projection effects resulting from the line of sight of the observation. Such possibilities can only be examined by properly modeling the 3D field of the solar eruption.

On the other hand, we expect Alfvénic vortex shedding to be largely independent of the external magnetic field, since its occurrence is solely due to the presence of a sharp obstacle with respect to the flow direction (convex regions in our simulation).

Forward modeling of the eruption revealed that the projection effects associated with optically thin observations of solar



**Figure 3.** (a) Sample field lines of the eruption (red). Green arrows indicate the direction of the velocity field in the  $xy$ -plane. (b) Isosurface of vorticity at  $|\hat{v}| = 1.75 \times 10^{-2} \text{ s}^{-1}$ , colored by  $\hat{v} \cdot \hat{\mathbf{B}}$ . Dashed line shows the region where  $\hat{v} \cdot \hat{\mathbf{B}} \approx 0$  (maximum KHI). Labels 1, 2, 3 indicate the approximate location of regions 1, 2, 3 of Figure 1(f). Black and gray arrows indicate the lines of sight used in Figure 4. Panels (c), (d) are side views of panel (b). Arrows indicate regions where  $\hat{v} \cdot \hat{\mathbf{B}} \approx \pm 1$  (minimum KHI) and regions of vortices associated with vortex shedding. (e) Side view of Figure 3(b), better for visualizing the vortex shedding regions. (f) Magnetic field strength at a vertical slice through the vortex shedding regions (also shown in panel (e)).



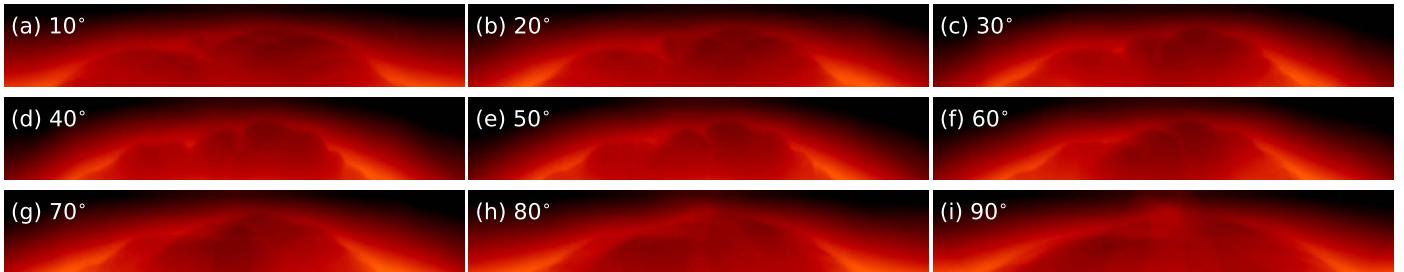
**Figure 4.** Synthetic observations of Fe IX for two lines of sight, one perpendicular to the plane of Figure 1 ( $x$ -axis,  $90^\circ$  polar angle, panels (a), (c), (d), (e), (f)) and another approximately perpendicular to the maximum KHI line of Figure 3(b) ( $45^\circ$  polar angle, panels (b), (g), (h), (i)). ((a), (b)) Line intensity showing the whole eruption. ((c)–(e)) Line intensity, Doppler velocity, and line width of region 1. (f) Line intensity of region 2 at  $90^\circ$  angle. ((g)–(i)) Line intensity, Doppler velocity, and line width of region 2 at  $45^\circ$  angle.

eruptions can significantly underestimate the presence of KHI in solar eruptions. This provides a simple explanation for the scarcity of the detection of KHI features in observations of solar eruptions. On the other hand, corrugated CME fronts produced by KHI as in our model may have easily been disregarded, interpreted simply by differential CME expansion or other effects. Spectral diagnostics could potentially provide better evidence of KHI-related features than imaging diagnostics.

The studied eruption has the energy and physical scale of a small-scale eruption. Its magnetic field strength is expected to be weaker than that of a larger eruption. Therefore, it is possible that KHI develops easier at the periphery of smaller and less energetic events (including jets; e.g., Li et al. 2018, 2019), at least while the erupting fields are still located close to the solar surface so that

they are observed by EUV imagers and spectrometers. The formation of the KHI in small-scale eruptions can be a very useful local and global diagnostic to probe the magnetic field and local plasma properties of these events. This is further so, given that the KHI is highly sensitive to the angle between the velocity shear and the local magnetic field, thereby providing information on the global magnetic field topology.

Similar eruptions to the studied one can be produced in higher energies in a multiscale manner following a power-law distribution (Syntelis et al. 2019). For larger-scale eruptions, the magnetic field could be strong enough so that KHI develops less frequently closer to the solar surface. However, depending on the properties of the eruption, KHI could also develop away from the lower solar atmosphere, during the eruption’s expansion and propagation



**Figure 5.** Line intensity of region 2 for lines of sight between  $10^\circ$  and  $90^\circ$ .

inside the solar wind (Páez et al. 2017), as the shearing between a CME and the solar wind could be appropriate for the formation of KHI (Manchester et al. 2005). Therefore, our results should be applicable to larger-scale eruptions showing KHI either closer to or farther away from the solar surface.

The presence of turbulence in the solar wind and the existence of a shock and an expanding CME front is often invoked to explain the acceleration of particles associated with these events (Bemporad & Mancuso 2011). The generated turbulence associated with the KHI and Alfvénic vortex shedding as well as the resulting corrugated CME front demonstrated in our model should therefore be considered for turbulence and particle acceleration studies.

This work was supported by computational time granted from the Greek Research & Technology Network (GRNET) in the National HPC facility—ARIS. P.A. acknowledges funding from his STFC Ernest Rutherford Fellowship (No. ST/R004285/1). P.S. acknowledges support by the ERC synergy grant “The Whole Sun.”

### ORCID iDs

P. Syntelis <https://orcid.org/0000-0002-6377-0243>  
P. Antolin <https://orcid.org/0000-0003-1529-4681>

### References

- Antolin, P., De Moortel, I., Van Doorsselaere, T., & Yokoyama, T. 2016, *ApJL*, **830**, L22  
 Antolin, P., & Van Doorsselaere, T. 2019, *FrP*, **7**, 85  
 Antolin, P., Yokoyama, T., & Van Doorsselaere, T. 2014, *ApJL*, **787**, L22  
 Arber, T., Longbottom, A., Gerrard, C., & Milne, A. 2001, *JCoPh*, **171**, 151  
 Balbus, S. A. 2003, *ARA&A*, **41**, 555  
 Bemporad, A., & Mancuso, S. 2011, *ApJL*, **739**, L64  
 Bodo, G., Massaglia, S., Ferrari, A., & Trussoni, E. 1994, *A&A*, **283**, 655  
 Bonet, J. A., Márquez, I., Sánchez Almeida, J., Cabello, I., & Domingo, V. 2008, *ApJL*, **687**, L131  
 Chandrasekhar, S. 1961, *Hydrodynamic and Hydromagnetic Stability* (Oxford: Clarendon)  
 Faganello, M., & Califano, F. 2017, *JPlPh*, **83**, 535830601  
 Fan, Y. 2001, *ApJL*, **554**, L111  
 Fang, X., Yuan, D., Xia, C., Van Doorsselaere, T., & Keppens, R. 2016, *ApJ*, **833**, 36  
 Feng, L., Inhester, B., & Gan, W. Q. 2013, *ApJ*, **774**, 141  
 Ferrari, A., Trussoni, E., & Zaninetti, L. 1981, *MNRAS*, **196**, 1051  
 Foullon, C., Verwichte, E., Nakariakov, V. M., Nykyri, K., & Farrugia, C. J. 2011, *ApJL*, **729**, L8  
 Foullon, C., Verwichte, E., Nykyri, K., Aschwanden, M. J., & Hannah, I. G. 2013, *ApJ*, **767**, 170  
 Gómez, D. O., DeLuca, E. E., & Miniinni, P. D. 2016, *ApJ*, **818**, 126  
 Gruszecki, M., Nakariakov, V. M., van Doorsselaere, T., & Arber, T. D. 2010, *PhRvL*, **105**, 055004  
 Hasegawa, H., Fujimoto, M., Phan, T. D., et al. 2004, *Natur*, **430**, 755  
 Heyvaerts, J., & Priest, E. R. 1983, *A&A*, **117**, 220  
 Hillier, A., Barker, A., Arregui, I., & Latter, H. 2019, *MNRAS*, **482**, 1143  
 Howson, T. A., De Moortel, I., & Antolin, P. 2017a, *A&A*, **602**, A74  
 Howson, T. A., De Moortel, I., & Antolin, P. 2017b, *A&A*, **607**, A77  
 Johansen, A., Henning, T., & Klahr, H. 2006, *ApJ*, **643**, 1219  
 Karampelas, K., Van Doorsselaere, T., & Antolin, P. 2017, *A&A*, **604**, A130  
 Kuridze, D., Zaqrashvili, T. V., Henriques, V., et al. 2016, *ApJ*, **830**, 133  
 Leroy, M., Ripperda, B., & Keppens, R. 2018, *JGRA Advance Online Publication*  
 Li, X., Zhang, J., Yang, S., & Hou, Y. 2019, *ApJ*, **875**, 52  
 Li, X., Zhang, J., Yang, S., Hou, Y., & Erdélyi, R. 2018, *NatSR*, **8**, 8136  
 Ma, X., Delamere, P., Otto, A., & Burkholder, B. 2017, *JGRA*, **122**, 10382  
 Manchester, W. B., I., Gombosi, T. I., De Zeeuw, D. L., et al. 2005, *ApJ*, **622**, 1225  
 Martínez-Gómez, D., Soler, R., & Terradas, J. 2015, *A&A*, **578**, A104  
 Masters, A., Achilleos, N., Kivelson, M. G., et al. 2010, *JGRA*, **115**, A07225  
 Miura, A., & Pritchett, P. L. 1982, *JGR*, **87**, 7431  
 Möstl, U. V., Temmer, M., & Veronig, A. M. 2013, *ApJL*, **766**, L12  
 Nakariakov, V. M., Aschwanden, M. J., & van Doorsselaere, T. 2009, *A&A*, **502**, 661  
 Nykyri, K., & Foullon, C. 2013, *GeoRL*, **40**, 4154  
 Nykyri, K., & Otto, A. 2004, *AnGeo*, **22**, 935  
 Ofman, L., & Thompson, B. J. 2011, *ApJL*, **734**, L11  
 Páez, A., Jatenco-Pereira, V., Falceta-Gonçalves, D., & Opher, M. 2017, *ApJ*, **851**, 112  
 Páez, A., Jatenco-Pereira, V., Falceta-Gonçalves, D., & Opher, M. 2019, *ApJ*, **879**, 122  
 Pagano, P., Reale, F., Orlando, S., & Peres, G. 2007, *A&A*, **464**, 753  
 Ryu, D., Jones, T. W., & Frank, A. 2000, *ApJ*, **545**, 475  
 Slavin, J. A., Anderson, B. J., Baker, D. N., et al. 2010, *Sci*, **329**, 665  
 Syntelis, P., Archontis, V., & Tsinganos, K. 2017, *ApJ*, **850**, 95  
 Syntelis, P., Archontis, V., & Tsinganos, K. 2019, *ApJ*, **876**, 61  
 Terradas, J., Andries, J., Goossens, M., et al. 2008, *ApJL*, **687**, L115  
 Tian, C., & Chen, Y. 2016, *ApJ*, **824**, 60  
 Tsinganos, K. C. 1980, *ApJ*, **239**, 746  
 Van Doorsselaere, T., Antolin, P., Yuan, D., Reznikova, V., & Magyar, N. 2016, *FrASS*, **3**, 4  
 Zaqrashvili, T. V., Vörös, Z., Narita, Y., & Bruno, R. 2014a, *ApJL*, **783**, L19  
 Zaqrashvili, T. V., Vörös, Z., & Zhelyazkov, I. 2014b, *A&A*, **561**, A62  
 Zaqrashvili, T. V., Zhelyazkov, I., & Ofman, L. 2015, *ApJ*, **813**, 123  
 Zhelyazkov, I., Zaqrashvili, T. V., & Chandra, R. 2015a, *A&A*, **574**, A55  
 Zhelyazkov, I., Zaqrashvili, T. V., Chandra, R., Srivastava, A. K., & Mishonov, T. 2015b, *AdSpR*, **56**, 2727