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RESEARCH ARTICLE

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Key Points:

- Radial diffusion is enhanced from the late initial storm phase and peaks in the late main phase
- Enhancements are due to external driving and internal sources of Ultra Low Frequency wave power
- Storm time magnetic field diffusion coefficients are underestimated by Ozeke et al. (2014)

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ULF Wave Driven Radial Diffusion During Geomagnetic Storms: A Statistical Analysis of Van Allen Probes Observations

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Abstract The impact of radial diffusion in storm time radiation belt dynamics is well-debated. In this study we quantify the changes and variability in radial diffusion coefficients during geomagnetic storms. A statistical analysis of Van Allen Probes data (2012–2019) is conducted to obtain measurements of the magnetic and electric power spectral densities for Ultra Low Frequency (ULF) waves, and corresponding radial diffusion coefficients. The results show global wave power enhancements occur during the storm main phase, and continue into the recovery phase. Local time asymmetries show sources of wave power are both external solar wind driving and internal sources from coupling with ring current ions and substorms. Wave power enhancements are also observed at low L values ($L < 4$). The accessibility of wave power to low L is attributed to a depression of the Alfvén continuum. The increased wave power drives enhancements in both the magnetic and electric field diffusion coefficients by more than an order of magnitude. Significant variability in diffusion coefficients is observed, with values ranging over several orders of magnitude. A comparison to the Kp parameterized empirical model of Ozeke et al. (2014) is conducted and indicates important differences during storm times. Although the electric field diffusion coefficient is relatively well described by the empirical model, the magnetic field diffusion coefficient is approximately ~ 10 times larger than predicted. We discuss how differences could be attributed to data set limitations and assumptions. Alternative storm-time radial diffusion coefficients are provided as a function of L^* and storm phase.

Plain Language Summary The radiation belts are populations of highly energetic particles that are situated in the Earth's space environment and trapped by our global geomagnetic field. During active periods, the radiation belts experience dramatic structural changes and intense energization, that can be hazardous to local spacecraft. Therefore, understanding how and why the radiation belts change during these periods is a key goal of the Space Weather community. We use direct observations of magnetic and electric field fluctuations from the Van Allen Probes spacecraft to explore how electromagnetic waves can scatter particles toward and away from the Earth, in a process called radial diffusion. The results show that radial diffusion is significantly enhanced during active conditions, and also that there is a huge range of variability between different active periods. Estimates are compared to an existing model and analysis shows that the model tends to underestimate radial diffusion during these active conditions. The results show that radial diffusion may be more important than previously thought and provides estimates of radial diffusion for use in modeling radiation belt dynamics.

1. Introduction

Earth's inner magnetosphere is host to a population of highly variable, highly dynamic, and highly energetic particles known as the Van Allen radiation belts (Li & Hudson, 2019; Van Allen et al., 1958, 1959). Of particular interest is the outer radiation belt population that typically occupies radial distances greater than $3-4 R_E$ and is host to extremely energetic MeV electrons. During geomagnetic storms, this population undergoes dramatic enhancements as well as rapid flux dropouts (e.g., Baker et al., 2004; Murphy et al., 2018; Turner et al., 2012). The MeV electron component of the outer radiation belt can cause problematic satellite

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anomalies in space due to electrostatic charging and discharging, and the occurrence of these anomalies increases with geomagnetic activity (e.g., Choi et al., 2011; Love et al., 2000). Therefore, it is a key outstanding goal of the Space Weather community to fully understand the dynamics and variability of the radiation belt population during geomagnetic storms (Fry, 2012; Schrijver et al., 2015).

However, the source and loss processes of radiation belt electrons is complex and multifaceted (Reeves, et al., 2003). Radially localized peaks in the distribution of the radiation belts can arise due to local wave-particle interactions (Boyd et al., 2018; Claudepierre et al., 2013; Horne & Thorne, 1998; Horne, 2007; Horne, Thorne, Glauert, et al., 2005; Horne, Thorne, Shprits, et al., 2005; Reeves, et al., 2013; Thorne et al., 2013) and flux dropouts at the outer boundary (Hudson et al., 2014; Turner et al., 2012). In the presence of strong radial phase space density gradients, electrons can then be subjected to significant radial redistribution that acts to smooth the localized peaks. This process is radial diffusion. Radial diffusion of electrons can act to both accelerate electrons as they are transported earthwards and decelerate electrons as they are transported radially outwards, due to the conservation of the first adiabatic invariant. Consequently, radial diffusion has been attributed to playing an important and sometimes key role in radiation belt energization (Mathie & Mann, 2000; Rostoker et al., 1998; Ukhorskiy et al., 2005), rapid electron flux dropouts (e.g., Shprits et al., 2006), and the “slot-penetrating” supply to the inner radiation belt (e.g., Loto’aniu et al., 2006; Zhao & Li, 2013). Understanding the dynamics of radial diffusion during geomagnetic storms is of considerable importance.

In a stationary magnetic field, radiation belt electron motion forms closed paths around the Earth, termed drift shells. However, when these electrons are subjected to random electromagnetic fluctuations occurring on timescales comparable to their drift period, violation of the third adiabatic invariant occurs, and the electron is scattered radially. Radial diffusion describes the average rate at which the trapped population is radially scattered to new drift shells due to a large number of small electromagnetic perturbations (Fälthamar, 1965; Kellogg, 1959; Parker, 1960). Previous work has established that broadband ULF waves are a key driver of radial diffusion in radiation belt electrons, owing to wave frequencies (~1–10 mHz) that are comparable to the electron drift frequency (e.g., Brito et al., 2020; Elkington et al., 2003).

1.1. Diffusion Coefficients

Significant efforts have quantified the magnitude of potential radial diffusion due to ULF wave power into radial diffusion coefficients. In particular, Fei et al. (2006) assumed that the diffusion coefficient for a dipole field, D_{LL} can be separated into magnetic and electric field components, D_{LL}^B and D_{LL}^E , respectively. The total diffusion is then the sum of the magnetic and electric field components, $D_{LL} = D_{LL}^B + D_{LL}^E$. However, as demonstrated by Lejosne (2019) the assumption made by Fei et al. (2006) that the magnetic and electric fields are independent and uncoupled is incorrect, and using the approach of Fei et al. (2006) can lead to an underestimation of the total radial diffusion coefficients by a factor of 2. We note that this discrepancy is minor compared to the large variability of values observed, which is demonstrated in Section 3.2. Although we opt to use the Fei et al. (2006) formalism in this analysis, the associated assumptions should be kept in mind when interpreting.

Using the Fei et al. (2006) approach, the magnetic and electric field components of radial diffusion can be expressed as

$$D_{LL}^B = \frac{M^2}{8q^2\gamma^2 B_E^2 R_E^4} L^4 \sum_m m^2 P_m^B(m\omega_d) \quad (1)$$

$$D_{LL}^E = \frac{1}{8B_E^2 R_E^2} L^6 \sum_m P_m^E(m\omega_d) \quad (2)$$

where B_E is the equatorial magnetic field strength at the surface of the Earth, R_E is the radius of the Earth, L is the L-shell, M is the first adiabatic invariant, q is the electron charge, γ is the relativistic

gamma factor, and m is the azimuthal wave number. Equations 1 and 2 show that the diffusion coefficients are proportional to the sum of the power associated with the magnetic and electric field perturbations, $P_m^B(m\omega_d)$ and $P_m^E(m\omega_d)$ respectively. The field perturbations have an azimuthal wave number and wave frequency, ω , that satisfies the drift resonance condition ($\omega - m\omega_d = 0$) with an electron with bounce-averaged drift frequency ω_d . Note that the magnetic field perturbations considered correspond to the compressional component (parallel to the background magnetic field) and the electric field perturbations to the azimuthal component (perpendicular to the background magnetic field and in the east-west direction).

As demonstrated by Ozeke et al. (2014), these descriptions of the diffusion coefficients can be simplified by assuming that the wave power at a given m value can be represented as a fraction of the total observed wave power, giving

$$D_{LL}^B = \frac{L^8 4\pi^2}{9 \times 8B_E^2} \langle P^B(L, f) f^2 \rangle \quad (3)$$

$$D_{LL}^E = \frac{L^6}{8B_E^2 R_E^2} \langle P^E(L, f) \rangle \quad (4)$$

where $P^B(L, f)$ and $P^E(L, f)$ are the power of field perturbations, which is dependent on the frequency, $f = 2\pi/\omega_d$ of the perturbations as well as the L value. The angled brackets in Equations 3 and 4 indicate that the average power is taken over all frequencies. The analytical expressions provided in Equations 3 and 4 allow for the estimation of diffusion coefficients directly from measurements of the power spectral density of magnetic and electric field perturbations, at a given time and location. Note that this approach assumes that the power spectral density value is representative of the average power spectral density along an electron drift path. Several studies, such as Ozeke et al. (2014) and Brautigam and Albert (2000), have developed empirical models parameterizing the diffusion coefficients by the Kp index. It should be noted that the parameterization is of the ULF wave power itself, since this is the only observed quantity remaining in Equations 3 and 4. The empirical models provide an easily accessible method to estimate the magnitude of radial diffusion in the inner magnetosphere for a given time, requiring only knowledge of the Kp index and without relying on suitable spacecraft coverage. The results of studies such as Ozeke et al. (2014) and Brautigam and Albert (2000) have informed our understanding of how radial diffusion varies with L value and with geomagnetic activity. Ozeke et al. (2014) show that the electric field component of radial diffusion is consistently greater than the magnetic field component ($D_{LL}^E > D_{LL}^B$) for $1 \leq L < 7$. In contrast, Brautigam and Albert (2000) observe that, although the electric field dominates at low L values, the magnetic field diffusion coefficient is dominant for approximately $L > 4$. In both the Ozeke et al. (2014) and Brautigam and Albert (2000) Kp parameterizations, both D_{LL}^E and D_{LL}^B increase with L value and both D_{LL}^E and D_{LL}^B increase with Kp index. This suggests that during geomagnetic storms where the value of the Kp index is increased, the magnitude of radial diffusion is enhanced.

The enhancement in radial diffusion coefficients during geomagnetic storms is further supported by case studies, where event-specific radial diffusion coefficients are estimated from observations of the magnetic and electric field perturbations. Case studies such as Li et al. (2014), Olifer, et al. (2019), Pokhotelov et al. (2016) and Schiller et al. (2017) demonstrate significant and important deviations from the Kp-parameterized models. For example, Olifer, et al. (2019) observe that during the March 2015 geomagnetic storm the magnetic field component, D_{LL}^B , was consistently underestimated and the electric field component, D_{LL}^E , was consistently overestimated by the empirical model of Ozeke et al. (2014). The magnitude of mis-estimation varied throughout the event, and at times the difference between empirically modeled values and event-specific diffusion coefficients was multiple orders of magnitude. In this study, we aim to identify how and why radial diffusion varies during geomagnetic storms through a statistical analysis of ULF wave fluctuations observed by the Van Allen Probes. The results determine the typical variations during geomagnetic storms and establish whether case studies of single storms are characteristic of typical storm time variations.

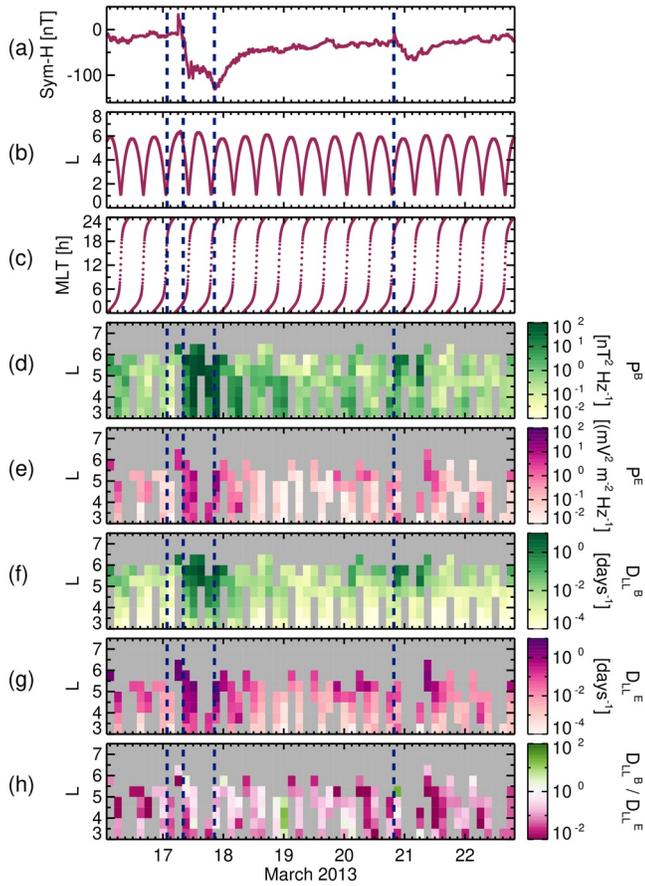


Figure 1. Timeseries of (a) Sym-H index [nT] and (b)–(h) Van Allen Probe A data during the March 2013 geomagnetic storm. The L value and MLT of the spacecraft is shown in panels (b) and (c), respectively. Panels (d)–(e) show data binned for time and L value, where the color of the bin indicates the median value of samples in the given bin. The time bin used has a width of 3 h and the L bin has a width of 0.5. The color scale is shown by the color bar on the right of the corresponding panel. The total power spectral density of the (d) magnetic field, P^B [$\text{nT}^2 \text{Hz}^{-1}$] and (e) electric field, P^E [$\text{mV}^2 \text{m}^{-2} \text{Hz}^{-1}$] is shown. The calculated diffusion coefficients for the (f) magnetic field, D_{LL}^B [days^{-1}] and (g) electric field, D_{LL}^E [days^{-1}] is shown. Panel (h) shows the median ratio of the magnetic field to electric field diffusion coefficients, D_{LL}^B / D_{LL}^E . Vertical dashed lines indicate storm phase timings, as identified using the (Walach & Grocott, 2019) algorithm. In chronological order, the dashed lines show the start of the initial phase, the start of the main phase, the start of the recovery phase, and the end of the recovery phase.

2. Data and Method

In situ observations of perturbations in the local magnetic and electric field were provided by the Van Allen Probes (Mauk et al., 2013). The Van Allen Probes consisted of two identically instrumented spacecraft (Probe A and Probe B) that commenced observations of the inner magnetosphere in October 2012, with operations ceasing in July 2019 for Probe B and October 2019 for Probe A. The spacecraft were in a 9 h orbit, with an apogee of $5.8 R_E$ geocentric radial distance, and an inclination of 10° . The orbital perigee was ~ 600 km during the prime phase, and began a period of perigee lowering maneuvers in January 2019. The orbital apogee precessed in local time, spanning all local times in less than 2 years. Due to the repeated sampling of the inner magnetosphere over a range of L and Magnetic Local Time (MLT) values for almost 7 years the Van Allen Probes present a highly suitable data set for statistical analyses of the outer radiation belt.

This study employs magnetic field observations from the Electric and Magnetic Field Instrument Suite and Integrated Science (EMFISIS) instrument (Kletzing et al., 2013) and electric field observations from the Electric Field and Waves (EFW) instrument (Wygant et al., 2013). The full data set includes observations from both Probe A and Probe B. We use Probe A data from January 01, 2013 to January 01, 2016, and Probe B from January 01, 2013 to June 16, 2017, where these time limits exclude degraded EFW data quality toward the end of the Van Allen Probes lifetimes. In the following subsection we detail how the magnetic and electric field observations are used to estimate radial diffusion coefficients.

2.1. Estimating Radial Diffusion Coefficients

Figure 1 shows an illustrative example of how the magnetic and electric radial diffusion coefficients are estimated for a typical geomagnetic storm that occurred during March 2013. Figure 1a shows the Sym-H index (Iyemori, 1990) trace during the event that exhibits the characteristic storm time variations. Following the approach of Walach and Grocott (2019), the Sym-H index was used to identify the event, as well as identifying the start and end times of the storm initial, main and recovery phases. The storm phase timings are marked by the vertical dashed lines in Figure 1. Further details about the storm identification are provided in Section 2.2. The L and MLT location of Van Allen Probe A is shown in Figures 1b and 1c.

The magnetic and electric field measurements were then taken for the full event, with a time resolution of 4 s. The spin axis electric field was suitably computed from the EFW observations under the assumption that $E \cdot B = 0$. In order to extract perturbations, the background magnetic and

electric field was first identified by taking a running average over a 20 min sliding window. The background field was subtracted and the residual field observations were transformed to a magnetic field-aligned coordinate system. The coordinate system was defined as having the parallel component aligned with the background magnetic field unit vector, the unit vector in the azimuthal direction is defined as oriented eastwards and perpendicular to the geocentric position vector, and the unit vector in the poloidal direction completes the coordinate system. Following the coordinate transformation, the parallel component of the magnetic field perturbation, B_{\parallel} , and the azimuthal component of the electric field perturbation, E_{ϕ} , was selected.

The perturbations were then used to estimate the associated power spectral densities using a Morlet wavelet transform. The power spectral density was limited to a frequency range of 1–15 mHz and an L range of

3–7.5. Lower L values were excluded as field perturbations cannot be reliably separated from the rapidly varying background field observed by the spacecraft at perigee. Figures 1d and 1e shows the total power spectral density (power summed over frequencies from 1 to 15 mHz) for the magnetic field and the electric field, where values have been binned and averaged for L value and time. The L value bin width is 0.5 and the time bin width is 3 h. For this storm, clear variations in power spectral density during the storm are apparent, with enhancements across all L values during the storm main phase. We note the presence of data gaps in the electric field power spectral densities. During periods of spacecraft charging and Earth eclipse events, electric field data is considered unreliable and omitted. This results in lower sampling of the electric field perturbations throughout the study, but we note that more than sufficient sampling is obtained for reliable statistical analysis.

The power spectral densities of the magnetic field, P^B , and electric field, P^E , were substituted into Equations 3 and 4, respectively, and provide estimates of the radial diffusion coefficients D_{LL}^B and D_{LL}^E . The estimation assumes that the observed power spectral densities are constant over all MLT values along the drift path, and the accuracy of this assumption is explored later. Note that throughout the manuscript, we explore the dependences and variations of the total power spectral density in the magnetic and electric field, P^B and P^E , to understand dynamics across the ULF wave band (e.g., Figures 1d and 1e). However, it is emphasized that the frequency-averaged wave power is used to calculate the corresponding radial diffusion coefficients, as mandated by Equations 3 and 4. Figures 1f and 1g shows the D_{LL}^B and D_{LL}^E values binned for L and time (in the same format as Figures 1d and 1e). As expected from Equations 3 and 4, the enhancements in power spectral density drive enhancements in both the D_{LL}^B and D_{LL}^E values during the storm main phase across all L values sampled. Studies, such as Olifer, et al. (2019), also identify important storm-time changes in the relative magnitudes of D_{LL}^B and D_{LL}^E , where D_{LL}^E is typically larger than D_{LL}^B prior to the storm and in agreement with empirical models (e.g., Ozeke et al., 2014). During the storm period, D_{LL}^B becomes comparable to, and at times dominates over, D_{LL}^E . Figure 1h shows the ratio of D_{LL}^B to D_{LL}^E and shows that generally for this event $D_{LL}^E > D_{LL}^B$, although the ratio approaches 1 during the recovery phase of the storm.

The analysis illustrated in Figure 1 was applied to all storms covered by the Van Allen Probes data set, providing a database of storm time P^B , P^E , D_{LL}^B and D_{LL}^E . Note that both Probe A and Probe B data was included. This database is statistically analyzed to explore how the power spectral density and radial diffusion coefficients vary on average during storm times, the variability across storms, and the physical drivers of the changes.

It is noted that although we have opted to use the method of Ozeke et al. (2014) to convert power spectral densities to diffusion coefficient values, other formalisms exist (see Lejosne and Kollmann (2020) for a thorough review). As mentioned previously, important differences can exist between the different approaches and it is highlighted that the approach adopted by Ozeke et al. (2014) (which follows work by Fei et al. (2006)) can lead to an underestimation of the total radial diffusion coefficients by a factor of 2 (Lejosne, 2019). Although it is shown in the subsequent analysis that this discrepancy is comparatively minor relative to the large variability in the observed values. Furthermore, a key aim of this study is to assess the reliability of the widely used empirical model of Ozeke et al. (2014), particularly focusing on the effect of parameterization by the Kp index. Employing a consistent method of calculation will aid in the comparison, and allow the impact of other factors (e.g., parameterization and datasets) to be isolated.

2.2. Identifying Geomagnetic Storms

In order to obtain the storm time database of radial diffusion coefficients, geomagnetic storms were identified for the Van Allen Probes data period. The storms were identified from timeseries of the Sym-H index using an automated algorithm developed and fully detailed by Walach and Grocott (2019). The Sym-H index was employed instead of the traditional Dst index (Sugiura & Poros, 1964) as it can be considered a high time resolution version of the Dst index (Wanliss & Showalter, 2006) and hence will provide higher accuracy of storm phase timings. Figure 1a shows a typical example of a storm identification using the Sym-H index. A geomagnetic storm can generally be separated into three distinct phases: The initial phase, the main phase, and the recovery phase. The start and end of these storm phases are indicated by the dashed vertical lines in Figure 1a. The initial phase is evident from a positive perturbation in the Sym-H index due

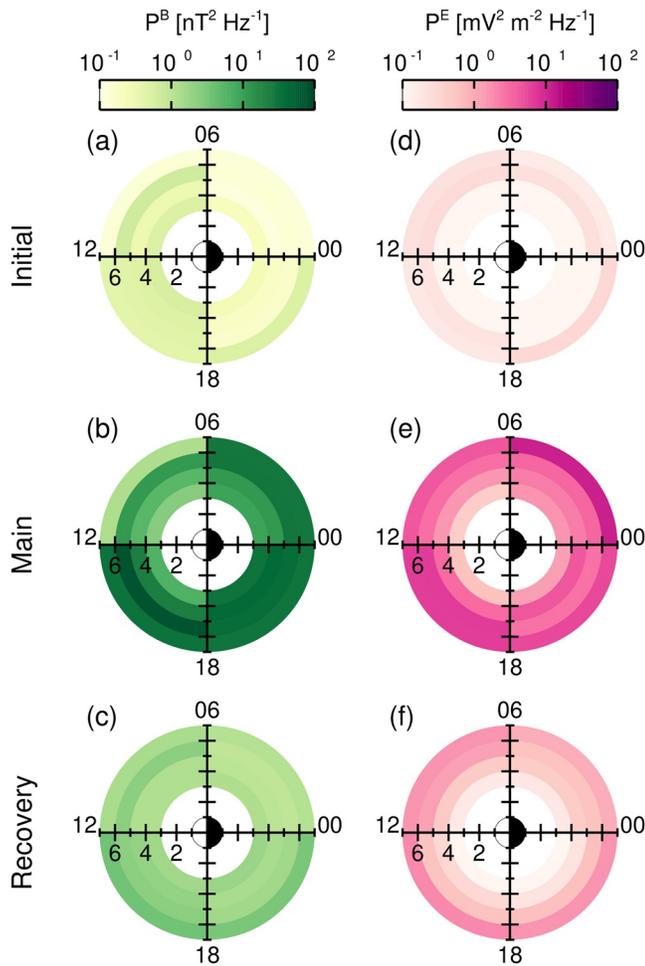


Figure 2. *L*-*MLT* maps where the bin color indicates median values of the (a–c) magnetic field power spectral density, P^B [$\text{nT}^2 \text{Hz}^{-1}$] and (d–f) electric field power spectral density, P^E , [$\text{mV}^2 \text{m}^{-2} \text{Hz}^{-1}$]. The power is summed over the 1–15 mHz frequency band. The *L*-*MLT* format maps show *L* value by radial location and *MLT* value by angular location of the bins. The data is also binned for storm phase as labeled.

to enhancements in the magnetopause currents and typically lasts ~ 1 day (Walach & Grocott, 2019). The initial phase is followed by the main phase, which is characterized by a rapid negative change in the Sym-H index due to significant and dramatic enhancements in the ring current population (e.g., Sandhu et al., 2021). The main phase has a typical duration of a few hours (Walach & Grocott, 2019). The final storm phase is the recovery phase, where the Sym-H index gradually recovers and increases to quiet time levels as the ring current decays (e.g., Sandhu et al., 2021). The recovery phase typically lasts several days (Walach & Grocott, 2019).

The Walach and Grocott (2019) algorithm first identifies a storm from any period where the Sym-H index crosses below a storm time threshold of -80 nT. The start of the recovery phase is marked as the time of the Sym-H index minimum. The start of the main phase and the end of the recovery phase are marked as the times immediately prior to and after the storm peak, where the Sym-H index is less than a quiet time level of -15 nT. The initial phase is defined to include the time of the Sym-H index maximum, and the start of the initial phase is marked as the time immediately prior to the Sym-H index maximum where the Sym-H index is at the quiet time level of -15 nT. Note that the storm time threshold (-80 nT) and the quiet time threshold (-15 nT) follow definitions by Hutchinson et al. (2011). The use of a -80 nT storm time threshold will naturally exclude any weaker storms that have a Sym-H index minima greater than this, and so the storms considered in this analysis will represent relatively large storm events. The -80 nT threshold was chosen to reliably avoid misidentification of other processes (such as large substorms, reconnection events, and magnetopause oscillations) that can manifest as weak depressions in the Sym-H index (Hutchinson et al., 2011). Therefore, the following analysis will reliably represent ULF wave activity and radial diffusion during strong geomagnetic storms. Future work will aim to explore the dependence of radial diffusion on storm intensity and storm drivers.

Overall, the algorithm identifies 45 storms during the Van Allen Probes data set. Over all these storms, the final data set comprises 8×10^6 samples of the magnetic field diffusion coefficient, and 5×10^6 samples of the electric field diffusion coefficient. This corresponds to total sampling time periods of 350 and 250 days, respectively.

3. Results

3.1. Power Spectral Density

First we explore spatial dependences in the ULF wave power, providing insight into the sources of the wave power and the key physical drivers of the radial diffusion coefficients. We show that important local time dependences exist, implying significant consequences for estimating event-specific radial diffusion coefficients from single spacecraft measurements.

Figure 2 shows spatial variations in the (a–c) P^B values and (d–f) P^E values, where power has been summed over the 1–15 mHz frequency band. Each row corresponds to a storm phase, as labeled, and shows samples binned for *L* and *MLT*. An *L* bin width of 1. And an *MLT* bin width of 6 h is used. The color of the bin indicates the median value of the parameter. The number of samples across the *L*-*MLT* bins is approximately uniform with each bin containing $\sim 10^5$ samples. Figure 2 highlights significant and storm-time dependent spatial variations in the power spectral densities of both the magnetic field, P^B , and the electric field, P^E . First focusing on the magnetic field component, Figure 2a shows that during the initial phase P^B is approximately 10^{-1} – 10^0 $\text{nT}^2 \text{Hz}^{-1}$, with slightly higher values observed on the dayside compared to nightside *MLT*

sectors. During the main phase, Figure 2b demonstrates that the P^B values are enhanced compared to the initial phase, with typical values of $\sim 10^2 \text{ nT}^2 \text{ Hz}^{-1}$. The enhancement is observed across nearly all spatial bins. The enhancement is smallest in the pre-noon MLT sector, imparting a weak local time asymmetry in the spatial distribution of P^B . Following the main phase, Figure 2c shows that the median P^B values subside to $\sim 10^0 \text{ nT}^2 \text{ Hz}^{-1}$ in the recovery phase. Values peak in the afternoon sector during both the main and recovery phases.

The power spectral density of the electric field also presents noteworthy storm-time variations. Figure 2d shows the median values of P^E in the initial phase are $\sim 10^{-1} \text{ mV}^2 \text{ m}^{-2} \text{ Hz}^{-1}$. No strong local time asymmetries are present in the initial phase. In contrast to the initial phase, Figure 2e shows that P^E is significantly enhanced during the storm main phase with values exceeding $\sim 10^1 \text{ mV}^2 \text{ m}^{-2} \text{ Hz}^{-1}$. Values are strongly enhanced across all MLT sectors, with an L dependence observed such that P^E tends to increase with L . Figure 2f shows that the values subside to $P^E \sim 10^0 \text{ mV}^2 \text{ m}^{-2} \text{ Hz}^{-1}$ in the recovery phase. A local time asymmetry is apparent with values peaking in the morning sector ($06 < \text{MLT} < 12$; Figure 2e).

Overall, Figure 2 shows that strong local time asymmetries arise during geomagnetic storms. Interestingly, the asymmetries and storm-time variations differ between the magnetic and electric fields. The physical processes that shape and contribute to the observed spatial dependences of P^B and P^E will be discussed in Section 4.

Although Figure 2 shows that there is important spatial structure in the power spectral density, the effect on a radiation belt electron is dependent on the power spectral density encountered by the electron on its drift path. In order to estimate the P^B and P^E values along an electron drift path, we bin data for L^* , where L^* is the third adiabatic invariant (Roederer, 1970; Roederer & Lejosne, 2018). The L^* value is calculated using the International Radiation Belt Environment Modeling (IRBEM) code (<https://sourceforge.net/projects/irbem/>) with the Tsyganenko and Sitnov (2005) magnetic field model for an equatorially trapped particle. It is noted that multiple methods for calculating L^* values exist, and this is a source of additional variability to the statistical analysis (Thompson et al., 2020). It is also stressed that in a distorted and non-dipolar magnetic field configuration, the L value and L^* values are not directly comparable. Asymmetries in P^B and P^E observed along the electron drift path are a combination of local time asymmetries displayed in Figure 2 and the asymmetry of the drift path.

Figure 3 shows the MLT variation of the median P^B and P^E , where each color profile corresponds to an L^* bin. The bars on each profile show the extent from the lower to upper quartile for each MLT bin. Figures 3a–3c shows profiles of the P^B values for the initial, main, and recovery phases, respectively. Figure 3a shows that during the main phase the power spectral density is quasi-symmetric in local time across all L^* values, with typical values of $\sim 10^{-1} \text{ nT}^2 \text{ Hz}^{-1}$. Values are slightly higher for dayside MLT sectors, but there is less than an order of magnitude variation across MLT sectors. During the main phase, the P^B values are enhanced across all MLT sectors. A local time asymmetry is prominent, with values peaking above $10^2 \text{ nT}^2 \text{ Hz}^{-1}$ in the post-noon and nightside MLT sectors (Figure 3b). In the recovery phase, Figure 3c shows that the local time asymmetry is subdued and P^B has reduced to $\sim 10^0 \text{ nT}^2 \text{ Hz}^{-1}$.

Figures 3d–3f shows the local time dependence of the P^E values. Similarly, to P^B , the P^E values are approximately symmetric in local time and have typical values of $\sim 10^{-1} \text{ mV}^2 \text{ m}^{-2} \text{ Hz}^{-1}$. The largest values are observed on the dayside. In the main phase the values typically do not exhibit any strong local time variations (Figure 2e). For the outermost L^* bin ($5 \geq L^* < 6$), the values are lowest in the morning sector. In the recovery phase, the observed P^E is reduced to approximately initial phase values.

In Figure 3 the outermost L^* values are not included as there are limited observations for $L^* > 6$, particularly during the main phase. During storm time conditions, magnetopause shadowing reduces the number of closed drift paths (e.g., Staples et al., 2020) and the last closed drift shell is often within $L^* = 7$ (Olifer, et al., 2018). Consequently, there are often no particles with closed drift paths in $6 \leq L^* < 7$ during geomagnetic storms. Across the observable L^* values, Figure 3 shows that there is limited variation with L^* and the profiles are typically within interquartile ranges.

An important consequence of the local time asymmetries is the spatial variations they impart to calculated diffusion coefficients. As radial diffusion is a drift averaged process, the radial diffusion coefficients should

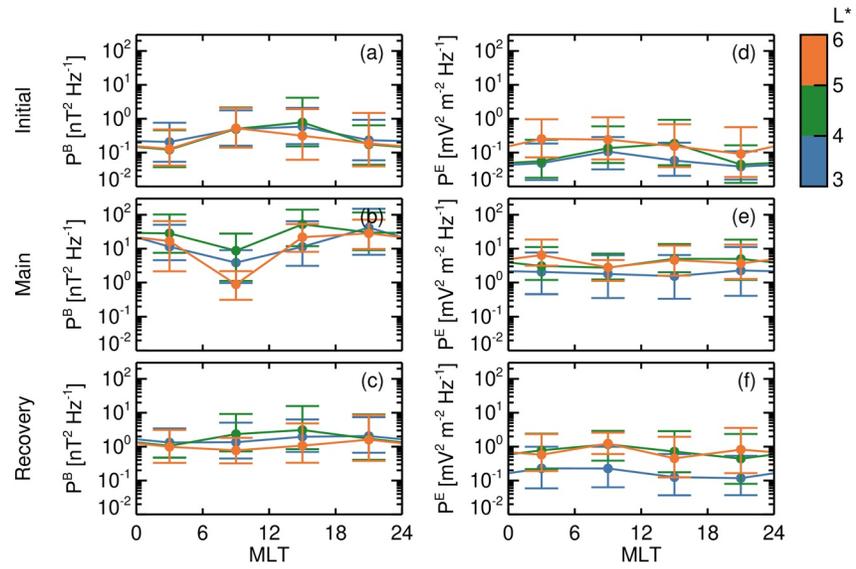


Figure 3. Median power spectral density as a function of MLT for (a–c) the magnetic field, P^B [$\text{nT}^2 \text{Hz}^{-1}$] and (d–f) the electric field, P^E , [$\text{mV}^2 \text{m}^{-2} \text{Hz}^{-1}$]. The data is binned for L^* , and each L^* bin is color coded according to the color bar shown on the right of the plot. The colored bars indicate the interquartile range for each MLT bin, using the same color coding as the corresponding profile. Data is also binned for storm phase, as labeled.

describe an average over all local times. However, Figures 2 and 3 show that during the main and recovery phases there are strong local time asymmetries observed in power spectral density and hence calculated radial diffusion coefficients. This implies that observations of the ULF waves that control radial diffusion during storm times are highly dependent on the local time position sampled. Figures 2 and 3 show that using spacecraft (single or multiple) or ground based magnetometers that sample a limited MLT sector for a single event can lead to significant under- or over-estimates of radial diffusion, as the approach would neglect spatial asymmetries. For example, event-specific radial diffusion coefficients estimated by Schiller et al. (2017) used Van Allen Probe observations in the pre-dawn sector. If ULF wave activity was significantly enhanced in, for example, the dusk sector, the estimated event-specific radial diffusion coefficients would be unrealistically low.

3.2. Radial Diffusion Coefficients

Section 3.1 demonstrates that the power spectral density of both the magnetic field and electric field varies along a drift path. As radial diffusion is a drift-averaged process, the following results bin data by L^* to describe the drift path and include contributions from all MLT sectors. Statistical estimates of radial diffusion coefficients are obtained by averaging over all MLT sectors for given L^* values.

Figure 4 shows the distribution of samples for (a) the magnetic field diffusion coefficients, (b) the electric field diffusion coefficients, and (c) the ratio of the magnetic field to the electric field diffusion coefficients. The distribution of the ratio is obtained by taking the ratio of D_{LL}^B to D_{LL}^E for each individual sample at a given time. The data have been further binned for storm phase, where the blue profiles correspond to the initial phase, the pink profiles correspond to the main phase and the orange profiles correspond to the recovery phase. Due to the comparatively longer duration of the recovery phase compared to the initial and main phases, the number of samples in the recovery phase is higher. The median value for each distribution is indicated by the position of the open colored circles above each panel, using the same color coding as the distributions. In order to reduce the spatial variability of the values, only samples within an L^* range of $4 \leq L^* < 5$ are shown here, and it is noted that the features are consistent for other L^* ranges. Figure 4 highlights some key features of the diffusion coefficients. First, the electric field diffusion coefficients are typically larger than the magnetic field counterparts (comparing Figures 4a and 4b), with the

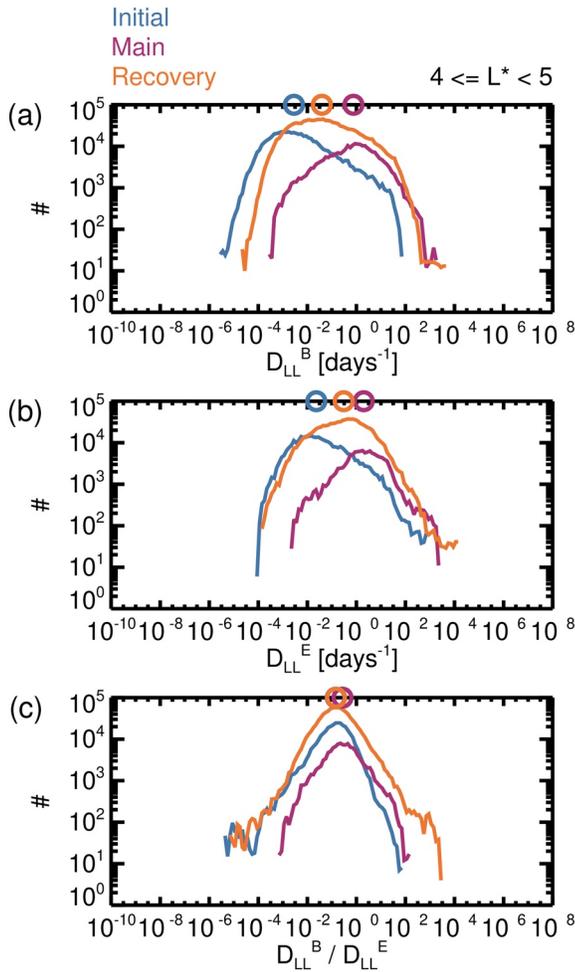


Figure 4. Occurrence distributions of (a) the magnetic field diffusion coefficient, D_{LL}^B , [days⁻¹], (b) the electric field diffusion coefficient, D_{LL}^E [days⁻¹], and (c) the ratio of the magnetic field to the electric field diffusion coefficients, D_{LL}^B / D_{LL}^E . Samples are within $4 \leq L^* < 5$. In each panel, the distributions are shown for samples taken in the storm initial phase (blue), main phase (pink) and recovery phase (orange). The median value of each distribution is indicated by an open circle at the top of the corresponding panel, using the same color coding as the distribution.

median values of D_{LL}^E approximately 10 times larger than D_{LL}^B . Figures 4a and 4b also show that all the distributions have a significant variability associated with them, with the distributions spanning several orders of magnitude for all storm phases. For both the magnetic field and electric field diffusion coefficients, the interquartile ranges of the distributions extend over approximately two orders of magnitude. A comparison of storm phases show that during the initial phase, the median diffusion coefficients are $D_{LL}^B = 3 \times 10^{-3}$ days⁻¹ and $D_{LL}^E = 2 \times 10^{-2}$ days⁻¹. During the main phase of the storm the values are considerably increased compared to the initial phase, with median values of $D_{LL}^B = 8 \times 10^{-1}$ days⁻¹ and $D_{LL}^E = 2 \times 10^0$ days⁻¹. In the storm recovery phase, the values return to lower values compared to the main phase, but remain slightly elevated compared to the initial phase distributions. The median values are $D_{LL}^B = 4 \times 10^{-2}$ days⁻¹ and $D_{LL}^E = 3 \times 10^{-1}$ days⁻¹. The differences in distributions between storm phases are all identified to be statistically significant according to the Kolmogorov-Smirnov test (with a p value threshold of 0.01), for both the magnetic field diffusion coefficients and the electric field coefficients.

Figure 4c shows the relative importance of D_{LL}^B and D_{LL}^E . As for the distributions shown in Figures 4a and 4b there is significant variability associated with the D_{LL}^B / D_{LL}^E distributions, with values spanning multiple orders of magnitude consistently for all storm phases. The median values of D_{LL}^B / D_{LL}^E are all located below 1, indicating that the electric field component of radial diffusion is typically dominant throughout the storm. This characteristic is consistent with empirical models (e.g., Ozeke et al., 2014). Although the median values are similar during the initial phase and the recovery phase ($D_{LL}^B / D_{LL}^E \sim 1 \times 10^{-1}$), the values are very slightly higher during the main phase ($D_{LL}^B / D_{LL}^E = 3 \times 10^{-1}$).

Although Figure 4 shows key features of the storm time diffusion coefficients, information regarding the temporal variability during a storm is lost by binning solely by storm phase. Alternatively, the storm time variations can be further investigated through a normalized superposed epoch analysis, and the results of this analysis are shown in Figure 5. As before, values within an L^* range of $4 \leq L^* < 5$ are shown to reduce variability due to spatial dependences and focus on storm time variations, although the variations are representative of the other L^* values. In Figure 5, a sample is binned according to its sample time relative to the duration of the current phase. For a given storm phase, the duration is scaled to the average duration of that phase across all storm events. These average durations are 37 h for the initial phase, 11 h for the main phase, and 61 h for the recovery phase. This method organizes the temporal variations based on the progression during a storm, and has been demonstrated to be an effective approach in organizing storm time behavior throughout the inner magnetosphere (e.g., Hutchinson et al., 2011; Murphy, et al., 2018; Wharton et al., 2020; Yokoyama & Kamide, 1997). Furthermore, this approach benefits by assigning each storm phase equal importance, as opposed to favoring a particular epoch in the non-normalized superposed epoch approach.

Figure 5a shows the Sym-H index for all storms in gray, and the median trace is shown by the blue profile using time bins of width 6 h. The vertical dashed lines demarcate the storm phases. It can be seen from Figure 5a that there is notable variability in the Sym-H index across storm events, particularly around the time of the Sym-H index minima. Figures 5b and 5c show contextual information on the solar wind conditions. Figure 5b shows the median solar wind speed, v_{SW} , as a function of time and Figure 5c shows the median southward component of the Interplanetary Magnetic Field (IMF), B_Z , as a function of time. The bars indicate the extent of the interquartile range for each time bin. The v_{SW} and B_Z parameters have been

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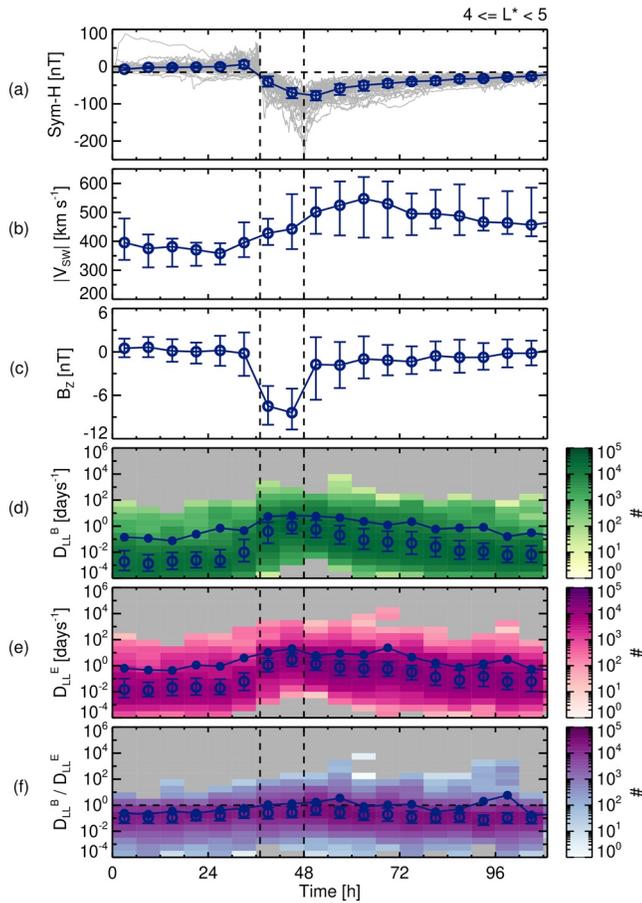


Figure 5. A normalized superposed multi-epoch analysis, where the dashed vertical lines show the epoch times used. Panel (a) shows the Sym-H index [nT] of all storms in gray, the median value for each time bin in blue, and the interquartile range is shown by the extent of the blue bars. The dashed horizontal line indicates the quiet time Sym-H value of -15 nT. Panels (b) and (c) show the median values of the solar wind speed magnitude, $|v_{sw}|$ [km s^{-1}] and the B_z component [nT] of the Interplanetary Magnetic Field (IMF), respectively. The interquartile range is indicated by the extent of the blue bars. Samples of the (d) magnetic field diffusion coefficient, D_{LL}^B , [days^{-1}], (e) electric field diffusion coefficient, D_{LL}^E [days^{-1}], and (f) ratio of the magnetic to electric field diffusion coefficient, D_{LL}^B / D_{LL}^E . Samples are within $4 \leq L^* < 5$. The samples are binned for time and the diffusion coefficient value and the number of samples in the bin is indicated by the bin color. The color scaling is indicated by the color bar on the right of each panel. The mean values for each time bin is shown by the solid blue profile. The median values are shown by the blue open circles and the interquartile ranges are shown by the blue bars. The dashed horizontal line in panel (f) indicates where $D_{LL}^B = D_{LL}^E$.

demonstrated to be dominant solar wind parameters in controlling ULF wave power (Bentley et al., 2018), and thus are expected to be important factors in shaping the storm time diffusion coefficients. Figure 5b shows that during the initial phase the solar wind speed is typically $350\text{--}400$ km s^{-1} . During the main phase the solar wind speed is progressively enhanced, peaking in the early recovery phase at $v_{sw} \sim 550$ km s^{-1} . Throughout the remainder of the recovery phase, v_{sw} remains elevated compared to the initial phase with median values of $\sim 400\text{--}500$ km s^{-1} , and large variability in values as indicated by the broad interquartile ranges. Similarly to v_{sw} , Figure 5c shows that the IMF B_z component is at low levels during the initial phase with values close to 0 nT. During the main phase, Figure 5c shows that B_z experiences a dramatic southward excursion and reaches values of ~ -8 nT by the late main phase. In the recovery phase the B_z values return to lower magnitudes with a slight southward bias. The interquartile ranges show significant variability throughout the early recovery phase.

Magnetic field diffusion coefficient, D_{LL}^B , samples are binned for time and the coefficient value, where the color of the bin in Figure 5d represents the number of samples in the bin. The blue closed circles show the mean value of the magnetic field diffusion coefficient for each time bin. The median and interquartile range is indicated by the blue open circles and the blue bars. Figure 5d demonstrates large variability throughout storms, with D_{LL}^B values spanning ~ 6 orders of magnitudes for any given time bin. This variability appears to be independent of storm phase, with similar interquartile ranges throughout the storm period. Figure 5d shows that D_{LL}^B is enhanced from the late initial phase and peaks in the main phase. The values then gradually subside throughout the recovery phase reaching values similar to the start of the storm.

The electric field diffusion coefficients, D_{LL}^E , are shown in Figure 5e, following the same format as Figure 5d. Similarly, to the magnetic field diffusion coefficients, the electric field counterpart exhibits large variability across the full storm duration with values typically spanning ~ 6 orders of magnitudes for any given time bin. D_{LL}^E experiences enhancements during the main phase of a storm. The D_{LL}^E values decrease toward original values throughout the recovery phase.

Figure 5f shows the ratio of the magnetic to electric field diffusion coefficients, D_{LL}^B / D_{LL}^E , using the same format as Figures 5d and 5e. The variability spans ~ 4 orders of magnitudes during the initial main phase, but is increased during the main and recovery phase. Throughout the initial phase, the values lie below 1, indicating $D_{LL}^B < D_{LL}^E$. The values begin to approach 1 during the main phase and D_{LL}^B and D_{LL}^E are roughly equivalent during the early recovery phase. Furthermore, Figure 5f shows there are an increased number of samples with $D_{LL}^B > D_{LL}^E$ in the late main and early recovery phase. In the late recovery phase, the values begin to subside and D_{LL}^E regains dominance.

Overall, Figures 4 and 5 show that values are enhanced from the late initial phase, continuing throughout the storm main phase and leading into the early recovery phase. Values subside to original levels by the end of the storm. A comparison to the Sym-H index and solar wind parameters show correlation with enhanced ring current conditions and solar wind driving. Figures 4 and 5 also suggest that the magnetic field diffusion coefficients experience larger enhancements, such that the magnetic and electric field counterparts are typically equivalent during the early recovery phase of a storm.

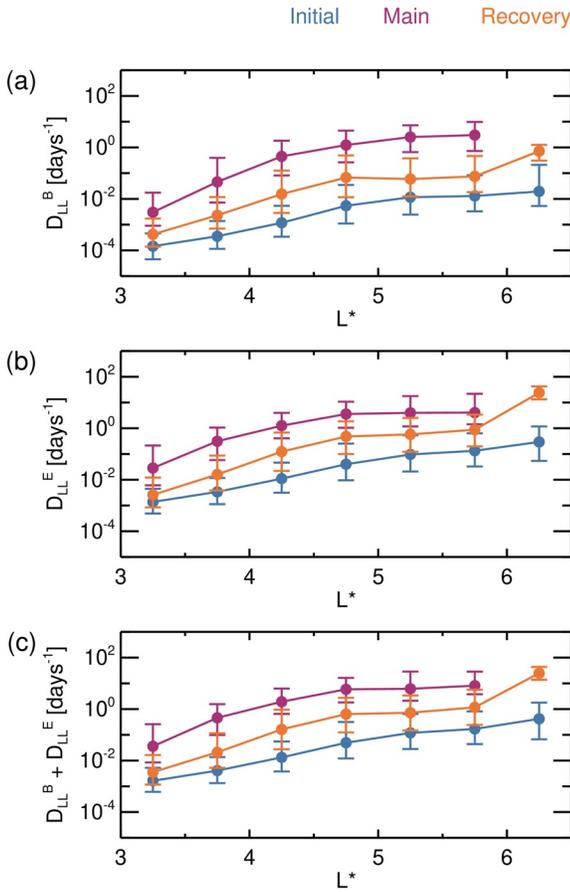


Figure 6. (a) Magnetic field diffusion coefficients, D_{LL}^B , [days⁻¹], (b) electric field diffusion coefficients, D_{LL}^E [days⁻¹], and (c) the sum of the magnetic and electric field diffusion coefficient, $D_{LL}^B + D_{LL}^E$ [days⁻¹] as a function of L^* . The profiles show the median values and the bars show the interquartile ranges. The profiles show data during the initial phase (blue), the main phase (pink) and the recovery phase (orange).

As Figures 4 and 5 are focused on the L^* range of $4 \leq L^* < 5$, Figure 6 investigates storm time changes across multiple L^* values. Figure 6a shows the median D_{LL}^B as a function of L^* , where the bars show the interquartile range of the samples within the bin. The blue, pink, and orange profiles correspond to samples in the initial phase, the main phase, and the recovery phase, respectively. Figure 6b shows the D_{LL}^E samples in the same format as Figure 6a. Figure 6 shows that both D_{LL}^B and D_{LL}^E increase with L^* , across all storm phases, with an approximate L^{*10} dependence. This is a consequence of both the diffusion coefficient formalism, where Equations 3 and 4 are shown to include dependences on L , as well as the slight radial gradients in P^B and P^E observed in Figures 2 and 3.

Figure 6 also demonstrates that the main phase enhancement in D_{LL}^B and D_{LL}^E is present across all L^* values. The magnitude of the main phase enhancement relative to the initial phase is approximately a factor of 10 for D_{LL}^B , and slightly less for D_{LL}^E . The recovery phase values are similar to the initial phase values across all L^* bins, with the medians often within the interquartile ranges.

As highlighted in the presentation of Figures 3 and 6 also shows that the higher L^* bins ($L^* \geq 5.5$) are often unsampled during the main phase of the storm. This is a direct consequence of the Earthwards displacement of the last closed drift shell during geomagnetic storms, as previously discussed.

Finally, Figure 6c shows the L^* profile for the sum of the magnetic and electric field diffusion coefficients, in the same format as Figures 6a and 6b. As expected, the profiles are similar to the electric field diffusion coefficients shown in Figure 6b, which dominate over the magnetic field contribution. The values shown in Figure 6c provide an estimate of the average diffusion coefficients specifically for storm time conditions, which may prove useful for the modeling and understanding of radial diffusion of typical geomagnetic storms.

3.3. A Comparison to Empirical Models

Widely used empirical models, such as Ozeke et al. (2014), are parameterized by the Kp index and show that during geomagnetically active conditions (high Kp values), both the magnetic and electric field diffusion coefficients are increased compared to quiet time conditions (low Kp values). In order to establish whether the Kp parameterized models are representative of typical storm time conditions, we directly compare the observed values to empirically modeled values using the Ozeke et al. (2014) model in Figure 7. The Ozeke et al. (2014) empirical model is given by

$$D_{LLmod}^B = 6.62 \times 10^{-13} L^{8.10} 10^{-0.0327L^2 + 0.625L - 0.0108Kp^2 + 0.499Kp} \quad (5)$$

$$D_{LLmod}^E = 2.16 \times 10^{-8} L^6 10^{0.217L + 0.461Kp} \quad (6)$$

For each sample of D_{LL}^B and D_{LL}^E , the corresponding empirical model values were calculated using Equations 5 and 6, where the L value and Kp value were taken at the time of measurement for each sample. The ratio of the observed diffusion coefficient to the modeled diffusion coefficient for the magnetic field (D_{LL}^B / D_{LLmod}^B) and the electric field (D_{LL}^E / D_{LLmod}^E) was taken for all samples. Figures 7a–7c shows the samples binned for D_{LL}^B / D_{LLmod}^B and L^* , where the color of each bin indicates the number of samples within the bins. The green profiles show the median D_{LL}^B / D_{LLmod}^B value for each L^* bin, where the interquartile ranges are shown by the green bars. The horizontal dashed line locates where the observed and modeled

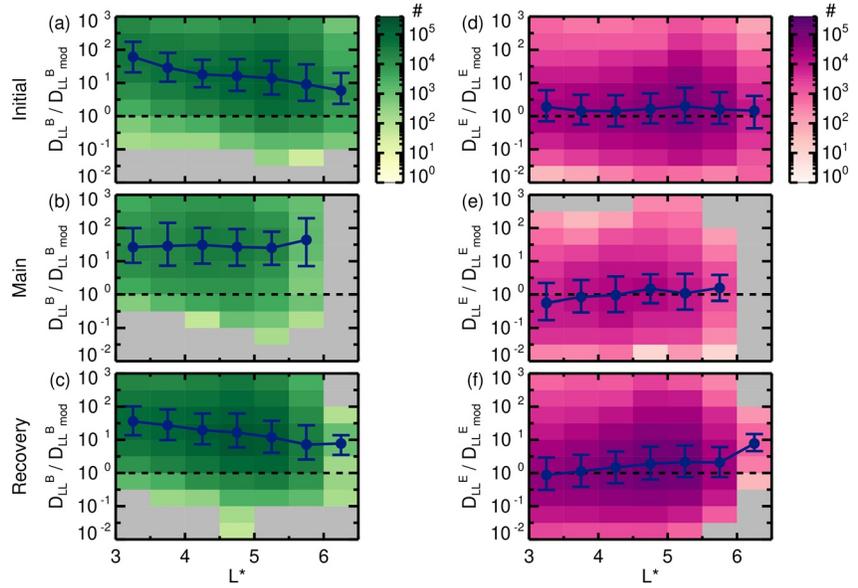


Figure 7. (a)–(c) Samples binned for the ratio of the observed to modeled magnetic field diffusion coefficients, D_{LL}^B / D_{LLmod}^B , and L^* . (d)–(f) Samples binned for the ratio of the observed to modeled electric field diffusion coefficients, D_{LL}^E / D_{LLmod}^E , and L^* . Modeled values were determined according to the Ozeke et al. (2014) model. The color of a given bin shows the number of samples within the bin, according to the color scale shown by the color bar on the right. The blue profiles show the median values as a function of L^* , and the blue bars show the corresponding interquartile range for each bin. Data is also binned for (a), (d) storm initial phase (b), (e) main phase, and (c), (f) recovery phase.

values are equal ($D_{LL}^B / D_{LLmod}^B = 1$). Each panel corresponds to a given storm phase, as labeled. For the magnetic field component, Figures 7a–7c shows that for all storm phases and L^* values, the values tend to lie above the $D_{LL}^B / D_{LLmod}^B = 1$ line and have typical values of 18. This indicates that the observed values are larger than the modeled values by a factor of ~ 10 during storm times. During the initial and recovery phases, Figures 7a and 7c shows an L^* dependence such that the difference is larger at low L^* values and smaller at high L^* values.

The corresponding analysis for the electric field diffusion coefficients is shown in Figures 7d–7f, in the same format as Figures 7a–7c. Values are binned for D_{LL}^E / D_{LLmod}^E and L^* . Similarly to the magnetic field component, the D_{LL}^E / D_{LLmod}^E values tend to lie above the $D_{LL}^E / D_{LLmod}^E = 1$ line during the initial and recovery phase, although the ratio has considerably smaller values of ~ 1.4 . It can also be seen that the values are much closer to 1 and lie slightly below the line at low L^* bins in the main phase (Figure 7e). Taking into account to variability of the data as shown by the interquartile ranges, the difference between observed and modeled values is not deemed to be significant for the electric field component.

Overall, the results show that although a Kp parameterized model does provide realistic storm time electric field diffusion coefficients, there is a tendency for the magnetic field diffusion coefficients to be underestimated by a factor of ~ 10 . The implications of this are discussed further in Section 4. It is also highlighted that the discrepancy between the observed and modeled values is comparable to the large degree of variability in values, as shown by Figure 4a and the interquartile ranges in Figure 6a.

4. Discussion

4.1. Why do Radial Diffusion Coefficients Increase During Storms?

The analysis presented in this manuscript demonstrates key features of the storm time radial diffusion coefficients. We find that both the magnetic and electric field diffusion coefficients exhibit strong and significant enhancements during geomagnetic storms across all L^* values observed. The enhancements are observed to

begin in the late initial phase of the storm, continue through the entirety of the main phase, and only begin to subside in the early recovery phase. The changes are driven by enhancements in the respective power spectral density, which are observed to have strong local time asymmetries (Figures 2 and 3). Here we will discuss potential and likely drivers of the power spectral density enhancements, and thus the enhancements in the radial diffusion coefficients.

It is well established that a key and dominant source of broadband ULF wave power in the inner magnetosphere is directly linked to external solar wind driving. Kelvin-Helmholtz instabilities at the magnetopause flanks generate ULF wave power locally (Chen & Hasegawa, 1974a, 1974b; Mann et al., 1999; Rae et al., 2005; Southwood, 1974) and buffeting of the magnetopause by solar wind pressure fluctuations enhances ULF wave power across the dayside (Kepko et al., 2002). Previous work supports these sources of ULF wave power by showing strong correlations between storm time wave power and solar wind pressure, IMF B_z , and coupling functions (Koskinen & Tanskanen, 2002; Posch et al., 2003; Takahashi et al., 2012). When only single solar wind parameters are considered, solar wind speed, IMF B_z , and variation in number density are the dominant parameters that account for a large degree of the variability in ULF wave power for all conditions (Bentley et al., 2018). For active conditions, the elevated solar wind speed associated with both CMEs and CIRs is observed to correlate with enhanced magnetospheric ULF wave power (Hudson et al., 2014; Simms et al., 2010; Zong et al., 2009). Due to the characteristics of the external wave power sources, local time asymmetries are imparted to the occurrence and amplitude of ULF wave activity such that they maximize in the dawn-noon MLT sector (Liu et al., 2009; Nosé et al., 1995; Nykyri, 2013; Pahud et al., 2009; Takahashi et al., 2016). The analysis presented in this manuscript shows that during the main phase and into the early recovery phase, where the solar wind driving displays dramatic enhancements (Figures 5b and 5c), corresponding enhancements in wave power are observed in the dawn-noon MLT sector and across the dayside (Figure 2). The wave power is elevated for both the magnetic and electric field components, and contributes to the corresponding enhancements in the diffusion coefficients.

However, Figure 2 also demonstrates that enhancements in ULF wave power occur globally during the main phase and recovery phase compared to the initial phase. Large enhancements are observed in the afternoon sector for the magnetic field fluctuations, and across the nightside sector for the electric field fluctuations. ULF wave power in these sectors is weakly linked to solar wind driving (Takahashi et al., 2012), and the sources of this wave power can instead be attributed to internal sources of ULF wave activity. Coupling of ULF waves and westward drifting ring current ions can drive enhancements in narrowband ULF wave activity predominantly in the afternoon and dusk MLT sectors (Baddeley et al., 2005; Engebretson & Cahill, 1981; Hughes, 1983; James et al., 2013, 2016; Nosé et al., 1998; Woch et al., 1990). Referring to Figure 2, afternoon sector enhancements during the main and recovery phase relative to the initial phase are apparent and are particularly prominent for the magnetic field component. The results suggest that ULF wave power excited by ring current ions contribute to the storm time radial diffusion coefficients, and could be dominant for the magnetic field component.

An additional and noteworthy internal source of ULF wave power is substorm activity. Studies, such as Nosé et al. (1998), show that substorms generate azimuthal ULF fluctuations across the nightside magnetosphere and peaks at 01–02 MLT. During the main phase of geomagnetic storms, substorm size and occurrence is markedly elevated and this enhancement in the source is reflected by main phase enhancements in ULF wave power for both the magnetic and electric field component across the nightside (Figures 2b and 2e). During the recovery phase the level of substorm activity subsides, and the postmidnight and broader nightside enhancements are reduced in magnitude.

Overall, storm time enhancements in ULF wave power occur across all MLT sectors due to a combination of both external and internal sources. These internal and external sources also exhibit important differences in the resulting azimuthal wave number, m . Studies show that externally driven waves that dominate the dayside MLT sectors typically have low m numbers (Rae et al., 2005; Tu et al., 2012), whereas ULF waves on the nightside have markedly higher m numbers (Barani et al., 2019; Fenrich et al., 1995; Sarris & Li, 2017). The high m ULF waves can originate from substorm-related activity, where they can be driven by drift-bounce resonances with injected ions (James et al., 2013; Walker et al., 1982; Yeoman et al., 2010). The presence of these high m number waves has implications for our estimations of radial diffusion coefficients. Whereas the approach considers $m = 1$ in the calculations of D_{LL}^B and D_{LL}^E , the assumption can potentially lead to

underestimates of the diffusion coefficients of more than an order of magnitude (Barani et al., 2019; Tu et al., 2012). This assumption is a focus for future work, where we will more accurately account for variations in the m number and thus refine the estimated storm time diffusion coefficients.

Another key feature of the storm time ULF wave power enhancements is associated with the radial distribution. Figures 2 and 3 show that the enhancements extend across all L and L^* values observed, and the wave power penetrates to low radial distances. During quiet and non-storm time conditions, external solar wind driving is a dominant source of ULF wave power and the ULF wave amplitudes are expected to decay exponentially from the source (i.e., the dayside magnetopause) (e.g., Southwood, 1974). Furthermore, wave power propagating from the dayside magnetopause can be effectively reflected at the plasmopause boundary (Lee et al., 2002). In contrast with quiet time conditions and in agreement with the observations presented here, previous studies show that during storm times enhanced wave power is observed at very low radial distances and within L values of ~ 4 (Georgiou et al., 2018; Lee et al., 2007; Loto'aniu et al., 2006; Rae et al., 2019). As highlighted by these studies, the accessibility of wave power to low L values can be enabled by changes in the Alfvén continuum. Studies, in particular Wharton et al. (2020), show that Eigen frequencies across a large range of L values are significantly depressed during the main phase of the storm. This is attributed to a weakening of the local magnetic field strength due to the magnetic field perturbations associated with an enhanced ring current population (Jorgensen et al., 2004; Sandhu et al., 2018), as well as an increase in the local plasma mass density due to enhanced concentrations of heavy ions (e.g., Kale et al., 2009; Kistler & Mouikis, 2016; Sandhu et al., 2017; Yue et al., 2019). Furthermore, studies such as Kale et al. (2009), Katus et al. (2015) show that during the main phase of the storm the plasmopause boundary moves to lower L values. This Earthward displacement can exceed several L values for extreme cases. The formation of plasmaspheric plumes in the afternoon sector during the main and recovery phase of the storm can also “trap” fast mode waves within the plumes (Degeling et al., 2018), which may explain the afternoon sector enhancement observed in Figures 2b and 2c. Overall, the suppressed Eigen frequency continuum in conjunction with the Earthwards displacement of the plasmopause allows ULF wave power to access lower L values during the main phase, as evidenced by the enhanced wave power at low L values (Figure 2). Power enhancements at low L values can also be enabled by the compression of the magnetopause, acting to move the external ULF wave power source Earthwards. Murphy et al. (2015) shows that for both CME and CIR driven storms, the compression of the magnetopause during the main phase independently contributes to significant enhancements in power at low L values.

The results demonstrate that both internal and external sources of wave power contribute significantly to enhanced radial diffusion coefficients during geomagnetic storms. It is also suggested that large scale changes in the shape and internal conditions of the magnetosphere result in enhanced diffusion coefficients at low radial distances.

4.2. How Reliable is the Kp Parameterized Model During Storms?

A key aim of the study was to identify whether a simple Kp parameterization, often used in empirical models such as Brautigam and Albert (2000) and Ozeke et al. (2014), provide an accurate description of radial diffusion coefficients during geomagnetic storms. As established by Figure 7, the magnetic field diffusion coefficients are systematically higher by ~ 10 times compared to the empirically modeled values. The electric field diffusion coefficients are also higher, but by a smaller multiplication factor of ~ 2 . This offset between observed and empirically modeled values is consistent across all L^* values and during all storm phases. It is also interesting to note the large spread of samples in Figure 7, which highlights that the discrepancy between observed and empirically modeled values can be highly variable for single storm samples.

Differences from the empirical model are supported by storm time case studies of radial diffusion. For example, Jaynes et al. (2018) and Olifer et al. (2019) demonstrated that during the main phase of the March 2015 storm the observed diffusion coefficients varied significantly compared to values using the Ozeke et al. (2014) model. The magnetic field diffusion coefficients were larger than the model values, whereas the electric field diffusion coefficients were lower than the model values. This is roughly consistent with the trends shown in Figures 7b and 7e. In contrast, a statistical study by Murphy et al. (2015) showed that observed main phase electric field diffusion coefficients were typically larger by more than an order of

magnitude compared to the values predicted by the Ozeke et al. (2014) model, and the overestimation was most prominent at low L values. However, Murphy et al. (2015) also note the large degree of variability observed in the diffusion coefficients, similarly to that observed in this study. Discrepancies with empirical models are further supported by Murphy et al. (2016), where it was shown that empirical Kp parameterized models can under/over-estimate electric field diffusion coefficients by four orders of magnitude similarly to the results shown in Figures 7d–7f.

The driver of discrepancies between observed and model values may arise through the measurement and analysis techniques. First, we consider differences in the magnetic field component. The Ozeke et al. (2014) empirical model for the magnetic field diffusion coefficient is based on observations from the Geostationary Operational Environmental Satellite (GOES) and the Time History of Events and Macroscale Interactions during Substorms (THEMIS) spacecraft. These observations cover L values from 5–7. Coverage of the non-geosynchronous L values is provided by the THEMIS observations, which are restricted to a time period of 2007–2011. The resulting empirical model will not account for the large enhancements in power at low L values, and the sampled L values are likely to lie outside of the last closed drift shell during storm times. It can be argued that the region assessed in this study lies outside the coverage of the empirical model leading to discrepancies in the magnetic field diffusion coefficient. In addition, the temporal coverage of the THEMIS spacecraft mostly corresponds to the solar minimum phase, whereas the Van Allen Probes measurements cover the solar maximum and declining phases. Studies such as Murphy et al. (2011) show that ULF wave power is enhanced during the solar maximum and declining phases, suggesting that the diffusion coefficients measured in this study will be larger than diffusion coefficients measured in the solar minimum phase. This supports the underestimation of the magnetic field diffusion coefficient by the Ozeke et al. (2014) empirical model (Figures 7a–7c). Finally, a key difference between the Ozeke et al. (2014) model and the results shown here are the choice of drift shell coordinate. The Ozeke et al. (2014) model averages over values binned for L , whereas we choose to average over L^* as it is a more realistic description of electron drift trajectories. The difference between L and L^* can be significant at large radial distances (approximately $L > 5$) (Roederer, 1967). However, as the radial dependence of power is small compared to, for example, local time asymmetry and event variability, the impact of using L as opposed to L^* is not expected to be the dominant source of differences.

The Ozeke et al. (2014) description of the electric field diffusion coefficients is based on ground magnetometer measurements from stations in the Canadian Array for Realtime Investigations of Magnetic Activity (CARISMA) (Mann et al., 2008) and U.K. Sub-Auroral Magnetometer Network (SAMNET) (Yeoman et al., 1990) arrays. The data is restricted to samples in the dayside sector (06–18 MLT) and covers L values from 2 to 8 (Ozeke et al., 2012). A dipole magnetic field is assumed for the mapping from station magnetic latitude to L value. Although the L coverage is well suited for a comparison with the Van Allen Probes observations, the limitation of the Ozeke et al. (2014) to dayside MLT sectors and the assumption that the observed power is independent of local time may be important. The empirical model would not account for any power enhancements in the electric field in the nightside sector, which can be important in the main and recovery phase (Figures 2e and 2f). This would result in the model underestimating the average power over a drift orbit, and consequently underestimating the electric field diffusion coefficient. The assumption of a dipole magnetic field configuration may also contribute to discrepancies. During the main phase of a geomagnetic storm, the enhanced ring current significantly distorts and inflates the magnetic field configuration in the inner magnetosphere and field lines with a given footprint latitude can map to larger L values than during non-storm periods (Ganushkina et al., 2010; Sandhu et al., 2018). The electric field power on the dayside tends to increase with L so the magnetometer measurements would be overestimating values. This is reflected in the main phase comparison shown in Figure 7e, where observed values tend to be slightly lower than the empirically modeled values for low L^* bins. As with the magnetic field component, the L value is also used for drift-averaging instead of the L^* values, but as only the dayside sector is considered by Ozeke et al. (2014) the impact of using the L value is not considered to be significant.

A key feature of the Ozeke et al. (2014) model is the choice of parameterization. As shown by Figure 5, the diffusion coefficients exhibit important variations within the storm. In particular, the average values peak in the late main phase. This is a feature that cannot be extracted from Kp parameterizations, which averages together samples from different storm phases, as well as averaging over both storm time and quiet time sam-

ples. Studies have shown that the storm time ULF wave power is significantly larger than non-storm time values for the same Kp value (Dimitrakoudis et al., 2015), resulting in observed values being larger than the modeled values (Figure 7). These discrepancies support efforts and suggestions to consider parameterization by solar wind driving and storm epoch to more reliably describe radial diffusion coefficients (Bentley et al., 2019; Dimitrakoudis et al., 2015; Lejosne, 2020). Furthermore, probabilistic modeling approaches to space weather effects (e.g., Morley et al., 2018) present a key avenue for more accurate descriptions of storm time radial diffusion coefficients, allowing for the high degree of observed variability in diffusion coefficients to be represented.

5. Conclusion

This study presents a statistical analysis of Van Allen Probes observations of ULF wave power and estimated radial diffusion coefficients during geomagnetic storms. Dependences on storm phase and L^* were explored, and the temporal variation within storms were assessed using a superposed epoch analysis. Crucially, and in contrast to many previous studies, the local time dependence of wave power was accounted for when calculating the diffusion coefficients. We summarize the key findings here:

1. Storm time diffusion coefficients vary significantly and by several orders of magnitude during geomagnetic storms, experiencing enhancements in the late initial phase and maximizing in the late main phase. The results shown here provide storm time specific estimates of radial diffusion coefficients as a function of storm phase and L^*
2. Analysis suggests that increases in radial diffusion coefficients is a consequence of both internal and external sources of ULF wave activity, and also that the inner magnetospheric conditions (ring current intensity and mass density distribution) play a key role in the propagation of wave power during geomagnetic storms
3. A comparison to the empirical model of Ozeke et al. (2014) highlights some key differences, most notably that the model tends to underestimate the magnetic field diffusion coefficient during storm times

Throughout the analysis we have highlighted the large variability in radial diffusion coefficients during storm times. The range of values for a given epoch time spans several orders of magnitude for both the magnetic and electric components. Future analysis aims to explore the key contributors to variability during storm times. This will include a consideration of storm magnitude on shaping the diffusion coefficient values. The role of storm drivers (e.g., CME and CIR drivers) is also suggested to play an important role (Kalliokoski et al., 2020; Kilpua et al., 2015; Simms et al., 2010), and these effects on the magnitude and temporal variation of diffusion coefficients will also be assessed.

Data Availability Statement

The Van Allen Probes data is publicly available online (<https://cdaweb.gsfc.nasa.gov/index.html/>). The solar wind data and Sym-H index data are publicly available online (<http://wdc.kugi.kyoto-u.ac.jp/index.html>).

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